

BASS. XLIX. Characterization of highly luminous and obscured AGNs: local X-ray and [Ne v] λ 3426 emission in comparison with the high-redshift Universe

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ABSTRACT

We present a detailed analysis of the most luminous and obscured Active Galactic Nuclei (AGNs) detected in the ultra-hard X-ray band (14–195 keV) by *Swift*/BAT. Our sample comprises 21 X-ray luminous ($\log L_X/\text{erg s}^{-1} > 44.6$, 2–10 keV) AGNs at $z < 0.6$, optically classified as Seyfert 1.9 and 2. Using *NuSTAR*, *XMM-Newton*, *Suzaku*, and *Chandra* data, we constrain AGN properties such as absorption column density N_H , photon index Γ , intrinsic L_X , covering factor, and iron K α equivalent width. We find median line-of-sight $\log N_H/\text{cm}^{-2} = 23.5^{+0.5}_{-1.2}$ and 2–10 keV rest-frame, de-absorbed

$\log L_X/\text{erg s}^{-1} = 44.7^{+0.8}_{-0.6}$, at the 5th and 95th percentiles. For sources with black hole mass estimates (12/20), we find a weak correlation between Γ and Eddington ratio (λ_{Edd}). Of these, six ($50 \pm 13\%$) lie in the $N_{\text{H}}\text{-}\lambda_{\text{Edd}}$ “forbidden region” and exhibit a combined higher prevalence of N_{H} variability and outflow signatures, suggesting a transitional phase where AGN feedback may be clearing the obscuring material. For the 13/21 sources with multi-epoch X-ray spectroscopy, $82^{+6}_{-16}\%$ exhibit variability in either 2–10 keV flux ($73^{+9}_{-16}\%$) or line-of-sight N_{H} ($33^{+15}_{-10}\%$). For the 20/21 sources with available near-UV/optical spectroscopy, we detect [Ne v] $\lambda 3426$ in 17 ($85^{+5}_{-11}\%$), confirming its reliability to probe AGN emission even in heavily obscured systems. When renormalized to the same [O III] $\lambda 5007$ peak flux as $z = 2\text{--}9$ narrow-line AGNs identified with JWST, our sample exhibits significantly stronger [Ne v] $\lambda 3426$ emission, suggesting that high-redshift obscured AGNs may be intrinsically weaker in [Ne v] $\lambda 3426$ or that [Ne v] $\lambda 3426$ is more challenging to detect in those environments. The sources presented in this work serve as a benchmark for high-redshift analogs, showing the potential of [Ne v] $\lambda 3426$ to reveal obscured AGNs and the need for future missions to expand X-ray studies into the high-redshift Universe.

Keywords: Active Galaxies — AGN — Obscuration — Surveys — X-Rays — JWST — Optical/UV

1. INTRODUCTION

Active Galactic Nuclei (AGNs) are among the most energetic objects in the Universe, powered by the accretion of matter onto supermassive black holes (SMBHs) at the centers of their host galaxies (Antonucci 1993; Urry & Padovani 1995). AGNs have been studied across multiple wavelengths, allowing for detailed investigations into how obscuration and extinction impact our understanding of their intrinsic properties (e.g., Padovani et al. 2017; Hickox & Alexander 2018; Caglar et al. 2020). However, AGNs that are both heavily obscured (with equivalent hydrogen column densities $N_{\text{H}} > 10^{23} \text{ cm}^{-2}$) and highly luminous ($L_X > 10^{44} \text{ erg s}^{-1}$, 2–10 keV) remain underrepresented in population studies and therefore still not well understood (e.g., Ananna et al. 2019; Ni et al. 2021; Peca et al. 2024b; LaMassa et al. 2024). This is due to the rarity of highly luminous sources and because of the challenges of detecting heavily obscured objects in the UV, optical, and soft X-ray ($< 10 \text{ keV}$) bands (e.g., Hickox & Alexander 2018; LaMassa et al. 2019a; Peca et al. 2023, 2024b). Identifying such AGNs is a critical step in understanding these objects, as they likely represent a crucial phase in galaxy evolution, often linked to rapid SMBH growth and feedback processes that can regulate or suppress star formation in their host galaxies (e.g., Hopkins et al. 2006, 2008; Treister et al. 2012). In this context, all-sky surveys in ultra-hard X-rays ($> 100 \text{ keV}$), such as those carried out by the Swift/Burst Alert Telescope (Baumgartner et al. 2013; Oh et al. 2018) and INTEGRAL (Bird et al. 2016; Krivonos et al. 2022), meet both requirements, providing a nearly unbiased census of AGNs by detecting even the most obscured sources across the entire sky (e.g., Ricci et al. 2015, 2017a; Ananna et al. 2022a).

The local Universe ($z < 0.6$) provides a unique opportunity to study these AGNs with great detail. Due to their proximity and relative brightness, we can obtain high signal-to-noise ratio (SNR) spectra, including X-rays from telescopes like *NuSTAR* (Harrison et al. 2013), *XMM-Newton* (Jansen et al. 2001), *Suzaku* (Mitsuda et al. 2007), and *Chandra* (Weisskopf et al. 2000). These instruments offer a powerful combination of broad energy coverage and high sensitivity, allowing us to penetrate the obscuring material and characterize the properties of these heavily obscured AGNs. High-quality (i.e., high-SNR) spectra with sufficient counts are critical for reliably constraining absorption, which, in turn, enables accurate estimates of intrinsic luminosity. When combined with the black hole mass (M_{BH}), this allows the estimate of the Eddington ratio ($\lambda_{\text{Edd}} = L_{\text{bol}}/L_{\text{Edd}}$), which represents a key quantity for understanding how these SMBHs evolve and interact with their host galaxy environment (e.g., Ananna et al. 2022a; Ricci et al. 2022a). Furthermore, detailed local studies provide essential benchmarks for interpreting the plethora of new high-redshift ($z > 2$) AGN discoveries with the James Webb Space Telescope (JWST), which have proven challenging to characterize at other wavelengths due to limited SNR. Robust local observations thus enable the testing and refinement of theoretical models of AGN evolution and feedback, while also guiding predictions and observational strategies for current and future facilities (e.g., Ananna et al. 2022a; Sobolewska et al. 2011; Fabian et al. 2017; Marchesi et al. 2020; Peca et al. 2024a; Boorman et al. 2024).

This work presents a detailed characterization of a sample of heavily obscured and highly luminous AGNs selected from the BAT AGN Spectroscopic Survey (BASS; Koss et al. 2017, 2022a). BASS is a large-

scale spectroscopic survey of AGNs detected by the *Swift*/BAT telescope in the ultra-hard X-rays (14–195 keV), that leverages the vast multiband data available across the full electromagnetic spectrum, including X-ray (e.g., Ricci et al. 2017a; Gupta et al. 2021; Marcotulli et al. 2022; Tortosa et al. 2023), optical-UV (e.g., Oh et al. 2022; Gupta et al. 2024), near-infrared (e.g., Lamperti et al. 2017; den Brok et al. 2022; Ricci et al. 2022b), mid-infrared (e.g., Ichikawa et al. 2017, 2019; Pfeifle et al. 2023), millimeter (e.g., Koss et al. 2021; Kawamuro et al. 2023; Ricci et al. 2023a), and radio (e.g., Baek et al. 2019; Smith et al. 2020b,a). By leveraging *NuSTAR*’s hard X-ray (3–79 keV) capabilities to study obscuration in AGNs (e.g., Marchesi et al. 2018; LaMassa et al. 2019b; Kammoun et al. 2020), complemented by the softer X-ray coverage of XMM-*Newton*, *Suzaku*, and *Chandra* down to ~ 0.5 keV, we investigate the physical properties of these sources, further supported by the multi-band datasets from BASS. Additionally, using optical spectroscopy and focusing on the [Ne V] $\lambda 3426$ emission line, we compare our results directly with those emerging from recent JWST-selected AGN samples at high redshift, showing potential differences between local obscured AGNs and distant populations. Finally, we utilize our best-fit spectral models to illustrate how detailed studies of obscured AGNs could be expanded with future X-ray observatories such as *AXIS* and *NewAthena*.

The paper is organized as follows. In Section §2, we describe the sample selection and the multi-wavelength data used in this work. Section §3 describes the X-ray data reduction process, while Section §4 outlines the spectral modeling strategy adopted. Results from the X-ray and optical analyses are presented in Sections §5 and §6, respectively. We discuss our findings in Section §7, while a summary is provided in Section §8. Throughout this paper, we assumed a Λ CDM cosmology with the fiducial parameters $H_0 = 70 \text{ km s}^{-1} \text{ Mpc}^{-1}$, $\Omega_m = 0.3$, and $\Omega_\Lambda = 0.7$. Errors are reported at the 90% confidence level or as the 5th and 95th percentiles if not stated otherwise. Uncertainties in proportions are computed using the binomial distribution (Cameron 2011).

2. DATA AND SAMPLE SELECTION

2.1. Sample selection

We selected our sample from the BASS DR2 catalog (Koss et al. 2022a,b), which is based on the *Swift*/BAT 70-month catalog (Baumgartner et al. 2013), and from the upcoming DR3 release (Koss et al, in prep.), which expands the BASS sample by including newly detected AGNs from the extended *Swift*/BAT 105-month catalog (Oh et al. 2018). We selected sources with absorption-

corrected 2–10 keV luminosity $\log L_X/\text{erg s}^{-1} > 44.6$. The luminosity values for this selection were derived from the de-absorbed fluxes reported by Ricci et al. (2017a), who performed a detailed X-ray spectral analysis of sources from the *Swift*/BAT 70-month catalog with redshifts obtained from BASS DR1 (Koss et al. 2017). For the additional sources in the 105-month catalog, we estimated luminosities based on the 14–195 keV fluxes. To do so, we converted these to the 2–10 keV band using a power-law model with a photon index of $\Gamma = 1.8$, which corresponds to the median value for non-beamed AGNs in Ricci et al. (2017a). This approach leverages the fact that the 14–195 keV flux is relatively unaffected by obscuration up to $N_H \sim 10^{24} \text{ cm}^{-2}$ (Ricci et al. 2015; Ananna et al. 2022a), ensuring consistency across catalogs while expanding the sample with the latest *Swift*/BAT data. To focus on obscured AGNs, we followed the optical classifications in Koss et al. (2022b) excluding Seyfert 1 objects. This resulted in selecting sources with a narrow H β emission line and either a broad (FWHM $> 1000 \text{ km s}^{-1}$) or narrow H α emission line, corresponding to Seyfert 1.9 and Seyfert 2, classifications, respectively. Additionally, we excluded beamed AGNs and removed one source (BAT ID 1240) with high Galactic extinction, $E(B-V)=0.99$, and Galactic $N_H \sim 7 \times 10^{21} \text{ cm}^{-2}$ (Kalberla et al. 2005), as this level of foreground extinction introduces significant uncertainty in the intrinsic optical classification, emission line measurements, and X-ray column density analysis. Our final sample consists of 21 sources at redshift $z < 0.6$, of which six are optically classified as Seyfert 1.9 and 15 as Seyfert 2 (see Figure 1). Of these 21, 12 are from the BASS DR2 catalog, while nine are from the upcoming DR3, effectively almost doubling the number of high-luminosity and optically obscured AGNs. We expect no additional sources meeting our criteria in the forthcoming BASS DR3 release, as optical spectroscopy has already been obtained for the relevant *Swift*/BAT 105-month sources.

2.2. X-ray data

All sources have been observed with *NuSTAR* at least once, and five of them twice (BAT IDs 32, 80, 199, 505, 714). To provide additional coverage at softer X-ray energies, we incorporated XMM-*Newton* data, where available, for 11 sources (BAT IDs 80, 119, 199, 476, 505, 714, 787, 1204, 1296, 1346, and 1515). Among these, two sources had multiple XMM-*Newton* observations (BAT IDs 80 and 1204). Since BAT ID 1204 has been observed several times, we selected two observations, prioritizing longer exposure times and minimizing the temporal separation between observations to ensure

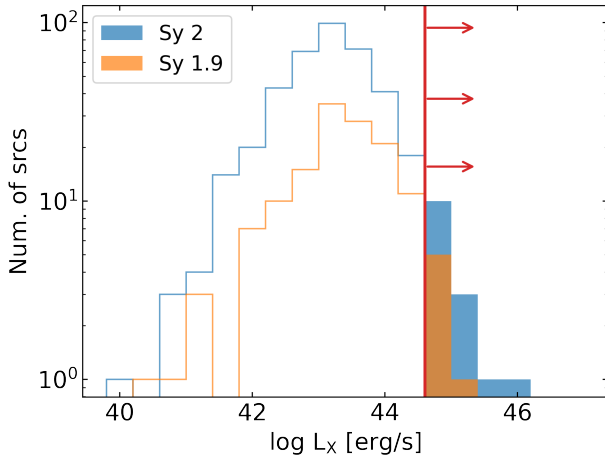


Figure 1. The selected sample consists of 21 AGNs with $\log L_X/\text{erg s}^{-1} > 44.6$, indicated by the red line, classified as Seyfert 1.9 and 2 (Koss et al. 2022b). The filled histograms represent the selected sources, while the full histograms show the full sample of Seyfert 1.9 and 2 (orange and blue, respectively) from BASS DR3 (Koss et al. in prep.). Additional details on the sample selection are provided in the text.

high-SNR spectra while mitigating potential variability-induced biases in the soft X-ray band (e.g., Laha et al. 2025). Given its large collecting area over a broad energy range (0.5–10 keV) and its greater availability for our sources compared to other soft X-ray telescopes, XMM-Newton was prioritized for soft X-ray coverage when available. Additionally, for two sources lacking XMM-Newton coverage (BAT ID 1248 and BAT ID 32), we included Chandra and Suzaku observations, respectively, to supplement the dataset. All the used X-ray observations are detailed in Table 1.

2.3. Optical data

Existing optical spectra were available for all 21 sources from the BASS survey: 15 from the Very Large Telescope (VLT) X-shooter instrument (Vernet et al. 2011), four from the Double Spectrograph (DBSP, Oke & Gunn 1982) on the 5m Hale telescope at Palomar Observatory, and two from the Sloan Digital Sky Survey (SDSS, York et al. 2000). Of these, 9/21 were unpublished in BASS DR2 Oh et al. (2022), and 1/21 (BAT ID 476) has been updated with a new VLT/X-shooter spectrum. For these 10 spectra, the fitting was performed following Oh et al. (2011, 2022) by de-redshifting the spectra and correcting them for Galactic foreground extinction, fitting the continuum emission using galaxy templates, and measuring the emission lines.

Additionally, for four sources (BAT IDs 555, 1248, 1296, and 1586), new optical integral field unit (IFU) observations were conducted with the Keck Cosmic Web

Imager (KCWI, Morrissey et al. 2018) on UT 30 March 2024 and UT 8 October 2024. KCWI was configured with the small slicer and the BL and RL gratings, centered at wavelengths of 4500 Å and 7150 Å, respectively, covering spectral ranges of 3420–5600 Å and 5600–8900 Å utilizing the 5600 Å dichroic with a total on-source observing time of 30 minutes. In this setup, the slitlets were 0.35'' wide in the east-west direction, with a pixel scale of 0.147''/pixel in the north-south direction and an 8.4'' × 20.4'' field of view. The average spectral resolution was $R \sim 3600$ for both gratings. Flux calibration was performed using nightly standards. Data reduction was completed with the standard KCWI DRP (v1.01). For five additional sources (BAT IDs 32, 199, and 1346), observations from the Multi-Unit Spectroscopic Explorer (MUSE) instrument (Bacon et al. 2010) at the VLT were used. The fully reduced phase 3 “MUSE-DEEP” observations from the ESO archive were used¹.

2.3.1. Black hole masses

In this work, 12/21 (~57%) black hole masses were available for our sample, all computed using the velocity dispersion method of Koss et al. (2022c). Of these 12 masses, five (BAT IDs 20, 80, 714, 787, and 1204) were obtained from the BASS DR2 catalog, while the remaining seven (BAT IDs 32, 199, 555, 1248, 1296, 1346, and 1586) were newly estimated with the same procedure using the KCWI and MUSE data. In brief, the continuum and the absorption features were fit using the penalized PiXel Fitting software (pPXF; Cappellari & Emsellem 2004) to measure the central velocity dispersion for the galaxy, using templates from VLT/X-shooter, with the prominent emission lines masked. Measurements of the Ca II H and K $\lambda 3969, 3934$ Mg I $b \lambda 5183, 5172, 5167$ region (3880–5550 Å) were performed, as the redshift of the sources causes the Ca II $\lambda 8498, 8542, 8662$ triplet to shift into the near-infrared (NIR). We refer the reader to Koss et al. (2022c) for further details.

We excluded one M_{BH} measurement (BAT ID 476) from BASS DR2, which reported an unusually low black hole mass due to an apparent broad H α component. New VLT/X-shooter observations, however, show that the H α line is actually narrow, with a strong blue wing originating from [N II] $\lambda 6548$, rather than a true broad component. Unfortunately, the updated spectrum did not permit a revised M_{BH} estimate. For this object, as well as the other sources in our sample lacking black hole masses, the available spectroscopic data did not meet the criteria outlined in Koss et al. (2022c), which include

¹ <https://www.eso.org/rm/api/v1/public/releaseDescriptions/102>

Source Name	BAT ID	Redshift	XMM			NuSTAR		
			Date	PN/M1/M2	Exp.(ks)	Date	FPMA/B	Exp.(ks)
2MASX J00343284-0424117	20	0.213	-	- - -	-	2022-07-21	✓✓	20.7/20.5
2MASX J01290761-6038423	80	0.203	2017-10-31	✓✓✓	21.5/27.1	2016-12-03	✓✓	21.1/20.9
			2021-06-17	✓××	10.3	2020-08-30	✓✓	25.0/24.9
2MASS J02162672+5125251	119	0.422	2006-01-24	✓××	8.8	2024-12-18	✓✓	22.7/22.4
LEDA 2816387	199	0.108	2017-11-20	✓✓✓	12.2/19.2	2016-05-06	✓✓	24.6/24.4
			-	- - -	-	2021-06-04	✓✓	26.6/26.5
CXO J095220.1-623234	476	0.252	2009-01-20	✓✓✓	12.2/20.3	2023-07-12	✓✓	27.7/27.5
SDSS J102103.08-023642.6	494	0.294	-	- - -	-	2023-05-09	✓✓	19.6/19.4
SDSS J103315.71+525217.8	505	0.140	2018-06-02	✓✓✓	23.6/28.0	2018-06-02	✓✓	30.6/30.4
			-	- - -	-	2020-06-24	✓✓	23.0/22.8
SDSS J113915.13+253557.9	555	0.219	-	- - -	-	2023-01-16	✓✓	21.1/20.8
LEDA 511869	714	0.076	2018-02-09	✓✓✓	6.3/17.3	2016-05-25	✓✓	23.6/23.4
			-	- - -	-	2020-03-25	✓✓	21.5/21.5
PKS 1549-79	787	0.150	2008-09-21	✓××	45.7	2016-07-12	✓✓	17.3/17.3
RBS 2043	1204	0.597	2013-11-13	✓✓✓	100.2/127.8	2021-01-07	✓✓	19.7/19.6
			2013-11-22	✓✓✓	80.5/92.1	-	- -	-
SWIFT J0231.4+3605	1241	0.321	-	- - -	-	2024-08-30	✓✓	20.7/20.5
2MASX J06215493-5214343	1291	0.209	-	- - -	-	2024-09-01	✓✓	19.6/19.4
GALEXASC J063634.15+591319.6	1296	0.204	2025-03-14	✓✓✓	15.7/19.7	2023-02-12	✓✓	22.7/22.4
2MASX J09003684+2053402	1346	0.235	2007-04-13	✓✓✓	9.9/15.7	2024-11-25	✓✓	22.7/22.5
2MASS J17422050-5146223	1515	0.218	2025-02-27	✓✓✓	12.1/17.2	2023-03-11	✓✓	16.1/16.0
SWIFT J2036.0-0028	1586	0.203	-	- - -	-	2022-10-13	✓✓	20.6/20.4
SWIFT J2103.3-2144	1595	0.289	-	- - -	-	2022-10-08	✓✓	19.1/18.9
APMUks(BJ) B233953.42-593202.5	1630	0.246	-	- - -	-	2022-06-22	✓✓	21.9/21.6

Source name	BAT ID	Redshift	Suzaku			NuSTAR		
			Date	XIS-FI/BI	Exp.(ks)	Date	FPMA/B	Exp.(ks)
ESP 39607	32	0.201	2010-12-19	✓✓	60.7/60.7	2023-05-08	✓✓	21.2/21.0
			-	- - -	-	2024-08-22	✓✓	21.6/21.4

Source name	BAT ID	Redshift	Chandra			NuSTAR		
			Date	ACIS-S	Exp.(ks)	Date	FPMA/B	Exp.(ks)
GALEXMSC J025952.92+245410.8	1248	0.206	2024-11-16	✓	9.9	2023-03-26	✓✓	22.3/22.1

Table 1. Data description for each source. From the left: Source name of the multi-band counterpart, BAT ID, redshift. The remaining columns provide the observation dates, camera availability, and exposure times (after data reduction) for XMM-Newton and NuSTAR (top), Suzaku and NuSTAR (middle), and Chandra and NuSTAR (bottom). Additional details on single sources are summarized in Appendix A.

adequate SNR and spectral resolution in the stellar absorption features. As a result, reliable M_{BH} estimates could not be obtained for nine sources in our sample. The new spectra and black hole masses will be included in the forthcoming BASS DR3 release.

3. DATA REDUCTION

3.1. NuSTAR

NuSTAR data were reduced using NuSTARDAS v2.1.2. We calibrated and cleaned the downloaded datasets using the standard `nupipeline` routine. The spectral extraction was done using the `nuproducts` com-

mand for the two FPMA and FPMB cameras. We defined circular extraction regions for each source centered on the source position for spectral extraction. Starting from 60'', roughly corresponding to the NuSTAR half-power diameter (HPD), we manually adapted the extraction radius to optimize the SNR (e.g., [Peca et al. 2021](#)). Considering all the sources, we used extraction radii in the range [35-60]''. The background was extracted from circular regions with a 90'' radius near each source, positioned on the same chip while avoiding contamination from nearby sources and CCD gaps. Each

background spectrum was inspected to ensure adequate sampling and to avoid empty channels.

3.2. *XMM-Newton*

XMM-Newton data were reduced using SAS v21.0.0, following the standard SAS threads for the EPIC PN, MOS1, and MOS2 cameras². In brief, after downloading the observation data files (ODFs), we produced the calibration index files and the summary files using the SAS commands `cifbuild` and `odfingest`. We reprocessed the ODFs using the commands `epproc` and `emproc` for PN and MOS cameras, respectively. Then, we filtered the events files for the flaring particle background, as described in the same SAS thread.

The spectral extraction procedure was similar to that of *NuSTAR* data, with source extraction radii starting at 30'' (corresponding to $\sim 90\%$ of the XMM encircle energy fraction) and then adapted for each source by optimizing the SNR. The final range of extraction radii was [25-40]''. The background was extracted from circular regions with different radii in the range [65-85]'', avoiding CCD gaps and nearby sources. Each background spectrum was checked to ensure good sampling and avoid empty channels. We ran `evselect` to extract the source and background spectra and calculated appropriate source and background scaling using the `backscal` command. Finally, the response matrices were obtained with `arfgen` and `rmfgen` SAS tasks.

3.3. *Suzaku*

In this work, we primarily used *XMM-Newton* to obtain coverage at energies softer than those observed with *NuSTAR*. For one source, ESP 39607 (BAT ID 32), we instead used *Suzaku* due to the unavailability of *XMM-Newton* observations. *Suzaku*/XIS data (Koyama et al. 2007) were used for this source. For most of its operational period, the XIS consisted of three cameras: the front-illuminated (FI) XIS 0 and XIS 3, and the back-illuminated (BI) XIS 1 (hereafter BI-XIS). Following Ricci et al. (2017a), we reprocessed the data for each XIS camera and extracted spectra from the cleaned event files using a circular aperture with a 1.7' radius centered on the source. The background was from a source-free annulus centered on the source, with inner and outer radii of 3.5 and 5.7', respectively. The extraction regions were selected following the same procedure described earlier. Response matrices were generated using the `xisrmfgen` and `xissimarfgen` tasks (Ishisaki et al. 2007). The spectra from XIS 0 and XIS 3 were

then merged using `mathpha`, `addrmf`, and `addarf` HEASoft³ tools.

3.4. *Chandra*

For one additional source, GALEXMSC 025952.92+245410.8 (BAT ID 1248), we used *Chandra* due to the unavailability of *XMM-Newton* observations. For the data reduction we used CIAO v4.16.0 (Fruscione et al. 2006) and followed the dedicated thread for point-like sources extraction⁴. First, we reprocessed the data using the `chandra_repro` tool. Then, we extracted the source and background spectra using the `specextract` command, where we specified the keywords `weight=no` and `correctpsf=yes` specific for point-like extraction. To determine source and background extraction regions, we followed the same procedure applied for *NuSTAR* and *XMM-Newton*. The source region was defined as a circle with a radius of 2.2'' while the background region was defined as an annulus with inner and outer radii of 8'' and 50'', respectively, avoiding CCD gaps and ensuring good background sampling.

3.5. Spectral binning

All the spectra were binned to a minimum of 20 counts/bin, as the number of net counts (i.e., background subtracted) exceeded 200 for all the spectra (e.g., Ricci et al. 2017a). For *NuSTAR* and *Chandra* the grouping was applied as part of the respective spectral extraction pipelines, for *XMM-Newton* we used the SAS tool `specgroup`, and for *Suzaku* we used the HEASoft `grppha` tool.

4. X-RAY SPECTRAL FITTING

We used PyXspec v2.1.3⁵ (Gordon & Arnaud 2021) to fit the extracted spectra. All available spectra were fitted simultaneously for each source, using the χ^2 statistic. We used Markov Chain Monte Carlo (MCMC) simulations for each fit, applying the Goodman-Weare algorithm with 50 walkers, a total chain length of 10^6 , and a burn-in phase of 10%. Each chain started after a first fit to assess the model parameters. The energy range used for the spectral fitting was 0.5–10 keV for *XMM-Newton*, 0.7–10 keV for *Suzaku*, and 0.5–7 keV for *Chandra*. For *NuSTAR*, the lower boundary was 3.0 keV, while the upper boundary was set in the range [30-50] keV depending on where the background starts to dominate the source emission.

³ <https://heasarc.gsfc.nasa.gov/docs/software/heasoft/>

⁴ <https://cxc.cfa.harvard.edu/ciao/threads/pointlike/>

⁵ Equivalent to XSPEC v12.14.0 (Arnaud 1996).

² <https://www.cosmos.esa.int/web/xmm-newton/sas-threads>

We fitted the sample using the following baseline model:

$$\text{Model} = \text{phabs} \times (\text{apec} + C_{\text{AGN}} \times (\text{Torus Model} + f_{\text{scatt}} \times \text{zcutoffpl})), \quad (1)$$

where *phabs* represents the Galactic absorption at the source position (Kalberla et al. 2005); *apec*⁶ is a thermal emission component from collisionally-ionized diffuse gas, modeling the soft X-ray emission from the host galaxy; and C_{AGN} allows the primary AGN emission to vary between spectra taken at different epochs. The *Torus Model* encapsulates absorption and reflection components from the obscuring material, as well as the primary cut-off power-law describing the intrinsic AGN continuum (see details in Section 4.1). We also included a secondary cut-off power-law component (*zcutoffpw*) to account for emission scattered into the line of sight and leaking through the torus. The photon index Γ of this secondary power-law was tied to that of the primary AGN emission, since it represents a fraction of the same intrinsic continuum, while its normalization was constrained to vary between 0 and 20% of the primary component through a multiplicative constant f_{scatt} , in line with typical scattering fractions found in previous AGN studies (e.g., Marchesi et al. 2016; Gupta et al. 2021; Peca et al. 2023). The cut-off energy of the power-law components is set to 200 keV, to match what is observed in the local Universe (e.g., Ricci et al. 2017a). In principle, an additional constant between the different instruments can be added to consider possible cross-calibration effects. However, we found this constant to be ~ 1.1 in agreement with other studies (e.g., Madsen et al. 2017; Baloković et al. 2021). Therefore, we did not include this additional parameter to avoid possible degeneracies (e.g., Marchesi et al. 2022).

For sources with only *NuSTAR* spectra available, the *apec* model and the secondary power-law component were turned off, as these components primarily contribute in the soft X-ray band not covered by *NuSTAR*, but accessible with *XMM-Newton*, *Chandra*, or *Suzaku*. A second *apec* component was added when needed to improve the fit results, as detailed in Section 5.1. Similarly, additional adjustments for individual sources were applied to improve the fit, as detailed in Appendix A.

The free parameters in our baseline model are the temperature and normalization of the *apec* component, the main power-law normalization, the photon-index Γ , the line-of-sight N_{H} , and the relative secondary power-law normalization, also known as the scattering frac-

tion, f_{scatt} . Additional free parameters from each torus model are included, as detailed in Section 4.1. Furthermore, for sources with multiple observations, we allowed C_{AGN} and the line-of-sight N_{H} to vary between different epochs. This allows for accounting for possible intrinsic flux variations and potential occultation events caused by the obscuring material. To avoid degeneracies between parameters, the photon index Γ was linked between the different epochs (e.g., Torres-Albà et al. 2023; Pizzetti et al. 2025).

4.1. Torus modelling

Several physically motivated torus models have been developed to study obscuration in AGNs, with many examples such as *borus02* (Baloković et al. 2018), *BNsphere* (Brightman & Nandra 2011), and *XCLUMPY* (Tanimoto et al. 2019, 2020). The need of multiple models arises from differences in their geometrical assumptions and treatment of obscuration, as highlighted in several studies (e.g., Saha et al. 2022; Kallová et al. 2024; Boorman et al. 2025). For this work, we selected the widely used *MYTorus* (Murphy & Yaqoob 2009a), *RXTorusD* (Ricci & Paltani 2023), which includes the interaction between X-ray photons and dust grains, and *UXCLUMPY* (Buchner et al. 2019), a clumpy torus model able to reproduce eclipsing events in the obscurer. A detailed description of these models is given below.

4.1.1. MYTorus

The *MYTorus* model (Murphy & Yaqoob 2009a) features a cylindrical, azimuthally symmetric torus with a fixed half-opening angle (i.e., the angular extent of the region where the torus does not obscure the central AGN) of 60° , filled with a uniform, neutral, cold reprocessing material. This model integrates the main components of an obscured AGN X-ray spectrum, such as obscuration, reflection, and fluorescent emission line components, using three distinct tables to treat them consistently: the zero-order continuum, the scattered component, and the emission lines component (See Eq. 2 and Eq 3).

First, we employed the model in its decoupled configuration (Yaqoob 2012; Yaqoob et al. 2015), which separates the line-of-sight column density ($N_{\text{H,los}}$) from the average column density ($N_{\text{H,ave}}$). In this configuration, the zeroth-order continuum (i.e., the photons that escape the torus without scattering) is independent of the inclination angle, fixed at $\theta_i = 90^\circ$. This makes the zeroth-order continuum a line-of-sight quantity unaffected by the torus geometry. We consider both edge-on and face-on scenarios to account for potential patchiness and varied torus configurations, as well as resulting Compton-scattering and emission line features. In the

⁶ Solar abundances were assumed (Anders & Grevesse 1989).

edge-on case, with $\theta_{i,S,L} = 90^\circ$, forward scattering is simulated and weighted by $A_{S,L90}$, representing a more uniform torus where photons are primarily reprocessed by the material between the AGN and the observer. In the face-on case, with $\theta_{i,S,L} = 0^\circ$, backward scattering is modeled and weighted by $A_{S,L0}$, indicating a patchier torus structure where photons scattered from the back of the torus have fewer interactions before reaching the observer. When $A_{S,L90}$ and $A_{S,L0}$ can vary freely, the configuration is called “decoupled”. In XSPEC syntax:

$$\begin{aligned} \text{Torus Model} = & \text{mytorusEzerov00.fits} \times \text{zcutoffpw} + \\ & A_{S,0} \times \text{mytorusscatteredH200v00.fits} + \\ & A_{L,0} \times \text{mytlV000010nEp000H200v00.fits} + \\ & A_{S,90} \times \text{mytorusscatteredH200v00.fits} + \\ & A_{L,90} \times \text{mytlV000010nEp000H200v00.fits}. \quad (2) \end{aligned}$$

where the Compton-scattered and emission line components are weighted differently using the multiplicative constants A_S and A_L , respectively. The additional free parameters when using this torus model are $A_{S,L,90}$, $A_{S,L,0}$, and the average $N_{H,\text{ave}}$. These parameters are linked between the spectra when multi-epoch spectra are fitted simultaneously.

Second, we utilized the MYTorus model in its “coupled” configuration, where the inclination angle, θ_i is allowed to vary freely:

$$\begin{aligned} \text{Torus Model} = & \text{mytorusEzerov00.fits} \times \text{zcutoffpw} + \\ & A_{S,\theta} \times \text{mytorusscatteredH200v00.fits} + \\ & A_{L,\theta} \times \text{mytlV000010nEp000H200v00.fits}. \quad (3) \end{aligned}$$

In this configuration, the scattered and emission line components represent the same inclination angle, and the only absorption parameter is the $N_{H,\text{los}}$, therefore reducing the number of model components. The additional free parameters of this torus model are θ_i , $A_{S,\theta}$, and $A_{L,\theta}$. As with the decoupled configuration, these additional parameters were tied across epochs when fitting multi-epoch spectra simultaneously.

4.1.2. RXTorusD

Next, we used the RXTorusD model (Ricci & Paltani 2023), which is based on the ray-tracing code for X-ray reprocessing, RefleX (Paltani & Ricci 2017). This model includes absorption and reflection from the torus with varying torus covering factors. In particular, the covering factor is defined as the ratio of the minor to the major axis of the torus (r/R). In this model, the $N_{H,\text{los}}$ is directly linked to the the equatorial column density

($N_{H,\text{eq}}$) via:

$$\begin{cases} N_{H,\text{los}}(\theta_i) = N_{H,\text{eq}} \left(1 - \frac{R^2}{r^2} \cos^2 \theta_i\right)^{1/2}, & \cos \theta_i < \frac{r}{R}, \\ N_{H,\text{los}}(\theta_i) = 0, & \cos \theta_i \geq \frac{r}{R}, \end{cases} \quad (4)$$

In its new version, RXTorusD considers the different effects of the interaction between X-ray photons and dust grains, such as dust scattering, near-edge X-ray absorption fine structures, and shielding. This model accounts for changes in the cross-sections of photon-gas interactions based on the fraction of metals in dust grains (the dust depletion factor). It also includes other physical processes, such as Rayleigh scattering and molecular gas scattering, which lead to significant differences in the predicted X-ray spectra for the same set of geometrical and physical parameters (Ricci & Paltani 2023).

We adopted the following torus configuration⁷, which accounts for reprocessed and continuum emission, respectively:

$$\begin{aligned} \text{Torus Model} = & \text{RXTorus_rprc.mod} + \\ & \text{RXTorus_cont.mod} \times \text{zcutoffpw}. \quad (5) \end{aligned}$$

The additional free parameters for this torus model component are the inclination angle θ_i and r/R , which are linked between the spectra when multi-epoch observations are fitted simultaneously.

4.1.3. UXCLUMPY

The last model we used was UXCLUMPY (Buchner et al. 2019). This model is designed to reproduce and model the column density and cloud eclipsing events in AGN tori, considering their angular sizes and frequency. UXCLUMPY differs from the previously mentioned models by including the torus clumpiness and cloud dispersion. The model aims to reproduce a cloud distribution with various hydrogen column densities based on observed eclipse event rates, assuming the clouds follow circular Keplerian orbits on random planes for simplicity. The distribution’s dispersion is controlled by the parameter $\text{TOR}\sigma$ ($\sigma \in [0 - 84]$), where a higher value indicates a larger dispersion and covering factor of the clouds. To model strong reflection features, UXCLUMPY includes an additional inner ring of Compton-thick ($\log N_{\text{H}}/\text{cm}^{-2} > 24$) material, whose covering factor is measured by the parameter CTKcover ($\in [0 - 0.6]$). $\text{TOR}\sigma$ and CTKcover together provide a robust framework to probe the torus geometry by modeling the cloud distribution and the extent of the reflecting material. In

⁷ From <https://www.astro.unige.ch/reflex/xspec-models>

the geometries discussed previously, $N_{\text{H,ave}}$ and $N_{\text{H,eq}}$ are closely related to $N_{\text{H,los}}$. In contrast, UXCLUMPY adopts a unified obscurer model, where a single clumpy torus geometry defined by TOR σ and CTKcover can be observed under a wide range of $N_{\text{H,los}}$ ($10^{20} - 10^{26} \text{ cm}^{-2}$). The used configuration⁸ is:

$$\text{Torus Model} = \text{uxclumpy-cutoff.fits} + f_{\text{scatt}} \times \text{uxclumpy-cutoff-omni.fits}. \quad (6)$$

the first table considers all the emissions from the primary component, and the second table accounts for the secondary power-law emission, therefore replacing the last term of Eq. 1. The additional free parameters of this torus component are the inclination angle θ_i , TOR σ , and CTKcover, which are linked between the spectra when multi-epoch spectra are fitted simultaneously.

5. RESULTS I: X-RAY CHARACTERIZATION

5.1. Determine the best-fit model

When fitting data with different models, selecting the best-fit model is a critical step, and several methods can be employed for this purpose (e.g., Buchner et al. 2014; Boorman et al. 2025). For instance, several works (e.g., Arcodia et al. 2018; Sicilian et al. 2022; Peca et al. 2023) emphasized the importance of using information criteria such as the Akaike Information Criterion (AIC; Akaike 1974) and the Bayesian Information Criterion (BIC; Schwarz 1978), which are applicable to both nested and non-nested models. These tests are likelihood-based, relying on the maximum likelihood obtained through chi-squared minimization, and they include a penalty term based on the number of free parameters to mitigate overfitting. While these criteria provide a reliable method for model selection, they do not consider the full parameter space of the likelihood distribution. In this work, we adopted the Deviance Information Criterion (DIC; Spiegelhalter et al. 2002) as our model selection metric. DIC is a Bayesian model selection approach that, like AIC and BIC, balances goodness of fit (from the likelihood function) and model complexity (the number of free parameters). However, DIC generalizes these criteria by incorporating the full posterior probability distribution, rather than relying on point estimates, to evaluate the best-fit model (Wilkins et al. 2022). To fully explore the posterior distributions required for DIC, we utilize the MCMC chains from the fitting procedure. The DIC is defined as:

$$\text{DIC} = \overline{D(\theta)} + p_{\text{DIC}}, \quad (7)$$

where $D(\theta) = -2 \log L(\theta)$ is the deviance, $\overline{D(\theta)}$ is the expectation of the deviance over the posterior distribution of the parameters, and p_{DIC} is the effective number of parameters. We used PyXspec to calculate the DIC, where p_{DIC} is determined according to the definition provided by Gelman et al. (2004),

$$p_{\text{DIC}} = \frac{1}{2} \overline{\text{var}(D(\theta))}; \quad (8)$$

i.e., half the variance of $D(\theta)$ over the chain. The model with the lowest DIC is preferred (Spiegelhalter et al. 2002; Gelman et al. 2004). The DIC criterion favors the UXCLUMPY model for 9/21 sources, the MYTorus (coupled) model for 5/21 sources, the RXTorusD model for 4/21 sources, and the MYTorus (decoupled) model for 3/21 sources. The best-fit spectra are shown in Appendix A.

In addition to this approach, we also tested simplified modeling where some of the geometrical parameters of the torus models were fixed. Specifically, we compared our baseline setup to a configuration where the inclination angle was fixed at 75° for MYTorus, with default values of $r/R=0.5$ for RXTorusD, and TOR $\sigma=28$ and CTKcover=28 for UXCLUMPY. For each pair of models (e.g., RXTorusD with and without free geometrical parameters), we evaluated the differences in DIC. A threshold of $\Delta \text{DIC} > 2$ was used, as it is considered to indicate “substantial” support for one model over another (e.g., Burnham & Anderson 2002). In all cases, the models with fixed parameters were not preferred ($\Delta \text{DIC} < 2$), supporting our choice to allow these geometrical parameters to vary, at least from a statistical perspective.

Table 2 summarizes the model selection and best-fit results. Parameters and associated uncertainties are derived from the MCMC chains and reported as the 50th percentile along with the 5th and 95th percentiles. All luminosities presented hereafter are intrinsic, meaning they are corrected for absorption and k-corrected to the rest-frame 2–10 keV band. For sources with multi-epoch observations, we report the average values of $N_{\text{H,los}}$ (hereafter referred to simply as N_{H}) and L_{X} , as our primary goal is to characterize the global properties of the sample. We verified that the N_{H} values from the different torus models are consistent within the uncertainties (see Appendix A). Single-epoch measurements are also provided and discussed in Section 5.6. The median values for the full sample are $\log N_{\text{H}}/\text{cm}^{-2} = 23.5^{+0.5}_{-1.2}$, $\log L_{\text{X}}/\text{erg s}^{-1} = 44.7^{+0.8}_{-0.6}$, and $\Gamma = 1.78^{+0.39}_{-0.17}$. For sources overlapping with the analysis of Ricci et al.

⁸ <https://github.com/JohannesBuchner/xars/blob/master/doc/uxclumpy.rst>

BAT ID	Redshift	Type	Flux (erg/s/cm ²)	Γ	log N _H (cm ⁻²)	log L _X (erg/s)	log λ_{Edd}
20	0.213	Sy2	27.2 ^{+0.8} _{-0.5}	1.71 ^{+0.21} _{-0.19}	23.7 ^{+0.5} _{-0.3}	45.0 ^{+0.3} _{-0.3}	-0.8 ^{+0.2} _{-0.2}
32	0.201	Sy2	14.7 ^{+0.5} _{-0.6}	1.72 ^{+0.11} _{-0.10}	23.6 ^{+0.2} _{-0.1}	44.7 ^{+0.1} _{-0.1}	-0.3 ^{+0.2} _{-0.2}
80	0.203	Sy2	4.9 ^{+0.2} _{-0.3}	1.64 ^{+0.23} _{-0.19}	24.0 ^{+0.2} _{-0.1}	44.4 ^{+0.3} _{-0.2}	-1.3 ^{+0.2} _{-0.2}
119	0.422	Sy2	56.9 ^{+3.3} _{-4.3}	1.72 ^{+0.06} _{-0.05}	22.5 ^{+0.2} _{-0.2}	45.5 ^{+0.1} _{-0.1}	...
199	0.108	Sy1.9	8.5 ^{+0.3} _{-0.4}	1.71 ^{+0.14} _{-0.12}	23.9 ^{+0.1} _{-0.1}	44.4 ^{+0.2} _{-0.2}	-0.4 ^{+0.2} _{-0.2}
476	0.252	Sy1.9*	41.5 ^{+0.8} _{-1.1}	1.85 ^{+0.07} _{-0.07}	23.0 ^{+0.1} _{-0.1}	45.0 ^{+0.1} _{-0.1}	...
494	0.294	Sy2	10.2 ^{+2.5} _{-5.4}	2.29 ^{+0.46} _{-0.40}	23.4 ^{+0.3} _{-0.4}	45.2 ^{+0.7} _{-0.6}	...
505	0.14	Sy2	3.4 ^{+0.2} _{-0.3}	1.85 ^{+0.13} _{-0.15}	24.1 ^{+0.6} _{-0.2}	44.1 ^{+0.1} _{-0.1}	...
555	0.219	Sy2	31.1 ^{+1.1} _{-1.7}	1.88 ^{+0.31} _{-0.20}	22.8 ^{+0.4} _{-0.5}	44.7 ^{+0.3} _{-0.3}	-0.9 ^{+0.2} _{-0.2}
714	0.076	Sy2	10.3 ^{+0.3} _{-0.5}	1.60 ^{+0.11} _{-0.10}	23.9 ^{+0.1} _{-0.1}	44.0 ^{+0.1} _{-0.1}	-1.5 ^{+0.2} _{-0.2}
787	0.15	Sy2	65.5 ^{+23.4} _{-28.8}	1.77 ^{+0.05} _{-0.03}	22.5 ^{+0.1} _{-0.1}	44.7 ^{+0.1} _{-0.1}	-0.4 ^{+0.2} _{-0.2}
1204	0.597	Sy2	58.4 ^{+0.5} _{-0.7}	1.72 ^{+0.08} _{-0.07}	23.5 ^{+0.1} _{-0.1}	45.7 ^{+0.1} _{-0.1}	-0.3 ^{+0.2} _{-0.2}
1241	0.321	Sy2	5.0 ^{+1.1} _{-2.5}	2.14 ^{+0.46} _{-0.41}	23.5 ^{+0.3} _{-0.5}	44.9 ^{+0.6} _{-0.7}	...
1248	0.206	Sy2	10.9 ^{+4.0} _{-5.3}	1.95 ^{+0.36} _{-0.26}	23.4 ^{+0.2} _{-0.1}	44.7 ^{+0.5} _{-0.4}	-0.2 ^{+0.2} _{-0.2}
1291	0.209	Sy2	5.3 ^{+1.1} _{-2.6}	1.81 ^{+0.55} _{-0.27}	23.5 ^{+0.3} _{-0.4}	44.5 ^{+0.8} _{-0.5}	...
1296	0.204	Sy1.9	22.0 ^{+0.5} _{-1.1}	1.60 ^{+0.08} _{-0.08}	22.6 ^{+0.2} _{-0.2}	44.6 ^{+0.1} _{-0.1}	-0.9 ^{+0.2} _{-0.3}
1346	0.235	Sy2	24.7 ^{+0.6} _{-0.5}	1.74 ^{+0.14} _{-0.14}	23.6 ^{+0.3} _{-0.1}	44.7 ^{+0.2} _{-0.2}	-1.1 ^{+0.2} _{-0.2}
1515	0.218	Sy1.9	27.2 ^{+1.6} _{-2.6}	1.72 ^{+0.03} _{-0.03}	21.6 ^{+0.5} _{-0.8}	44.6 ^{+0.1} _{-0.6}	...
1586	0.203	Sy1.9	23.0 ^{+0.4} _{-1.2}	1.86 ^{+0.15} _{-0.13}	23.2 ^{+0.5} _{-0.2}	44.7 ^{+0.2} _{-0.2}	-0.3 ^{+0.2} _{-0.2}
1595	0.289	Sy1.9	19.7 ^{+1.9} _{-2.0}	1.87 ^{+0.19} _{-0.15}	22.3 ^{+0.4} _{-0.8}	44.9 ^{+0.8} _{-0.2}	...
1630	0.246	Sy2	4.3 ^{+0.7} _{-1.8}	2.16 ^{+0.55} _{-0.36}	23.7 ^{+0.3} _{-0.3}	44.8 ^{+0.9} _{-0.5}	...

BAT ID	θ_i (deg)	f _{scatt} %	C.F.	kT (keV)	kT (keV)	Best Model	χ^2/dof
20	75 ⁺¹⁰ ₋₁₈	...	0.5 ^{+0.4} _{-0.2}	...	...	RXTorusD	149.8/172
32	70 ⁺¹⁶ ₋₈	0.9 ^{+0.3} _{-0.2}	...	0.4 ^{+0.1} _{-0.1}	...	MYTorus-coup.	479.9/473
80	69 ⁺¹⁷ ₋₇	0.2 ^{+0.2} _{-0.1}	...	0.8 ^{+0.7} _{-0.5}	...	MYTorus-coup.	195.2/191
119	70 ⁺¹⁴ ₋₁₉	0.7 ^{+0.4} _{-0.4}	0.8 ^{+0.2} _{-0.2}	...	...	RXTorusD	536.9/565
199	69 ⁺⁸ ₋₇	0.7 ^{+0.4} _{-0.3}	0.4 ^{+0.3} _{-0.2}	0.2 ^{+0.1} _{-0.1}	0.9 ^{+0.2} _{-0.1}	RXTorusD	343.5/314
476	...	0.6 ^{+0.3} _{-0.4}	...	0.6 ^{+2.2} _{-0.6}	...	MYTorus-dec.	613.8/639
494	35 ⁺³⁹ ₋₂₆	...	> 0.21	...	...	UXCLUMPY	36.9/31
505	...	0.5 ^{+0.3} _{-0.2}	...	0.6 ^{+1.0} _{-0.5}	...	MYTorus-dec.	107.9/118
555	...	...	...	...	...	MYTorus-dec.	133.4/140
714	> 72	0.3 ^{+0.1} _{-0.1}	...	...	...	MYTorus-coup.	251.1/233
787	> 77	2.3 ^{+2.6} _{-0.7}	0.4 ^{+0.1} _{-0.1}	...	...	UXCLUMPY	1187.2/1209
1204	> 83	14.4 ^{+2.5} _{-6.8}	...	1.9 ^{+5.0} _{-0.9}	7.1 ^{+2.4} _{-2.6}	MYTorus-coup.	7433.6/7019
1241	42 ⁺³⁴ ₋₃₁	...	0.6 ^{+0.4} _{-0.2}	...	...	UXCLUMPY	21.2/18
1248	42 ⁺³² ₋₃₀	< 17.6	0.7 ^{+0.3} _{-0.4}	...	...	UXCLUMPY	62.0/70
1291	43 ⁺³³ ₋₃₁	...	0.6 ^{+0.2} _{-0.4}	...	...	UXCLUMPY	32.7/28
1296	53 ⁺²³ ₋₃₁	7.9 ^{+5.5} _{-5.0}	0.5 ^{+0.2} _{-0.2}	...	...	UXCLUMPY	352.9/327
1346	77 ⁺⁹ ₋₁₂	3.7 ^{+8.4} _{-3.7}	...	1.3 ^{+6.1} _{-0.1}	3.1 ^{+0.1} _{-0.5}	MYTorus-coup.	958.3/952
1515	64 ⁺¹² ₋₃₆	...	0.5 ^{+0.1} _{-0.1}	...	...	UXCLUMPY	1054.3/1009
1586	76 ⁺¹⁰ ₋₃₆	...	0.4 ^{+0.3} _{-0.3}	...	...	RXTorusD	111.2/118
1595	58 ⁺³⁰ ₋₂₃	...	0.5 ^{+0.2} _{-0.3}	...	...	UXCLUMPY	74.3/81
1630	48 ⁺³¹ ₋₃₄	...	0.6 ^{+0.1} _{-0.4}	...	...	UXCLUMPY	18.3/20

Table 2. Results from the spectral analysis and derived quantities. *Top table*, from the left: BAT ID, redshift and optical classification from Koss et al. (2022b) (*: for BAT ID 476 see Section 2.3.1), observed 2–10 keV flux in units of 10⁻¹³ erg/s/cm², photon index Γ , logarithm of the line-of-sight N_H, logarithm of the intrinsic (rest-frame and absorption corrected) 2–10 keV luminosity, and logarithm of the Eddington ratio λ_{Edd} . *Bottom table*, from the left: Inclination angle θ_i , percentage of scattering fraction f_{scatt}%, covering factor C.F., temperature of the *apex* components, chosen best-fit model, and best-fit chi-squared over degrees of freedom. Note that the inclination angle for the MYTorus model in decoupled configuration is not reported, as it is not a parameter of the model, nor is it the covering factor for both MYTorus models (see text for details).

(2017a), we find good agreement, as detailed in Appendix B.

Out of the initial selection, we report that five sources are slightly below the luminosity threshold used for the sample selection, but still in the high luminosity regime ($L_X > 10^{44}$ erg/s). This minor discrepancy is likely due to using different datasets and models for the spectral analysis, as well as possible variability, given that the BAT data represent an average flux since *Swift*'s launch in 2004. We also find that one source (BAT ID 1515) is not classified as X-ray obscured ($\log N_H/\text{cm}^{-2} < 22$); however, this is consistent with the fact that sources classified as Seyfert 1.9 might indeed exhibit low levels of X-ray obscuration (e.g., Panessa & Bassani 2002; Oh et al. 2022). This corresponds to an X-ray obscured fraction for high-luminosity, Seyfert 1.9 and 2 sources of $95^{+2}_{-9}\%$ in our sample, consistent with the $85^{+4}_{-7}\%$ ($91^{+3}_{-16}\%$) from Ricci et al. (2017a) for the same optical selection and $\log L_X/\text{erg s}^{-1} > 44$ (> 44.6).

5.2. Photon index versus Eddington ratio

We investigate the relationship between the Eddington ratio (λ_{Edd}) and the X-ray photon index (Γ), which has been extensively studied by several works (e.g., Shemmer et al. 2006, 2008; Risaliti et al. 2009; Brightman et al. 2013, 2016; Trakhtenbrot et al. 2017; Liu et al. 2021; Huang et al. 2020; Kamraj et al. 2022), using samples selected in different ways. For instance, Risaliti et al. (2009) analyzed a sample of optically selected SDSS AGNs up to redshift 4.5, while Brightman et al. (2013) focused on unobscured AGNs in the COSMOS field up to redshift 2.5. This was later extended by Brightman et al. (2016), who studied megamaser Compton-thick AGNs in the local Universe using *NuSTAR*, and by Trakhtenbrot et al. (2017), who analyzed the full BASS DR1 sample leveraging the hard X-ray selection of *Swift*/BAT. More recent studies have further explored this relation for low- and high-Eddington regimes, with Huang et al. (2020) and Liu et al. (2021) finding steeper slopes or unobscured AGNs in the high-Eddington range at relatively low redshifts ($\lesssim 0.5$). In contrast, Kamraj et al. (2022) analyzed unobscured AGNs including *NuSTAR* data, and found no strong evidence of a significant correlation. These mixed results suggest that the $\lambda_{\text{Edd}}-\Gamma$ relation may not be universal and could depend on factors such as sample selection, luminosity range, X-ray energy coverage, spectral modeling techniques, and the methods used to determine λ_{Edd} .

Two primary explanations have been proposed to account for this correlation. The first links the observed trend to the impact of accretion disk radiation on the

AGN corona (e.g., Shemmer et al. 2008; Trakhtenbrot et al. 2017). At higher λ_{Edd} , intense UV/optical emission from the accretion disk enhances the cooling of the corona via Comptonization, lowering its temperature and optical depth, and producing steeper X-ray spectra. Conversely, at lower λ_{Edd} , the reduced photon supply results in a hotter and more compact corona, leading to flatter X-ray spectra. The second, alternative explanation, was proposed by Ricci et al. (2018), who showed through simulations that runaway pair production could drive this trend. They suggested that AGNs avoid the region in the temperature-compactness parameter space where pair production becomes significant, resulting in the temperature of the Comptonized plasma decreasing with the Eddington ratio. This effect can therefore explain the $\lambda_{\text{Edd}}-\Gamma$ trend, as the photon index is related to the temperature of the plasma (e.g., Ricci et al. 2018; Laha et al. 2025).

We computed the Eddington ratio, $\lambda_{\text{Edd}} = L_{\text{bol}}/L_{\text{Edd}}$, where L_{bol} is the bolometric luminosity and L_{Edd} is the Eddington luminosity, given by $L_{\text{Edd}} = 1.26 \times 10^{38} \times M_{\text{BH}}/M_{\odot}$ erg/s. L_{bol} was computed from the X-ray luminosities derived in this work using the Duras et al. (2020) bolometric correction. The median bolometric luminosity is $\log L_{\text{bol}}/\text{erg s}^{-1} \sim 46.1$. 12/21 sources in our sample have estimated M_{BH} , and therefore λ_{Edd} . The uncertainties in λ_{Edd} were estimated by combining in quadrature the error on L_X and the typical scatter in M_{BH} , which is dominated by systematics rather than uncertainties in emission line fitting (Koss et al. 2022c). For BASS DR2, the estimated uncertainty on M_{BH} is ~ 0.5 dex (Ricci et al. 2022b; Caglar et al. 2023), reflecting the intrinsic scatter in virial and velocity dispersion-based methods. The median value across our sample is $\log \lambda_{\text{Edd}} = -0.59$.

We found a weak correlation between λ_{Edd} and Γ . Using the `scipy.stats` module, we computed Spearman, Pearson, and Kendall's tau correlation coefficients of 0.59, 0.60, and 0.55, with corresponding p-values of 0.043, 0.036, and 0.014, respectively. These values correspond to a significance level between 2 and 3σ , which we consider indicative of a weak correlation. We performed a linear fit using the MCMC tool `emcee` (Foreman-Mackey et al. 2013), explicitly accounting for asymmetric uncertainties in both x and y. To do so, we carried out 1000 Monte Carlo realizations, in which the data points were randomly drawn from two half-Gaussian distributions centered on the best-fit value, with widths defined by the respective upper and lower uncertainties⁹.

⁹ <https://github.com/alessandropeca/BALinFit>

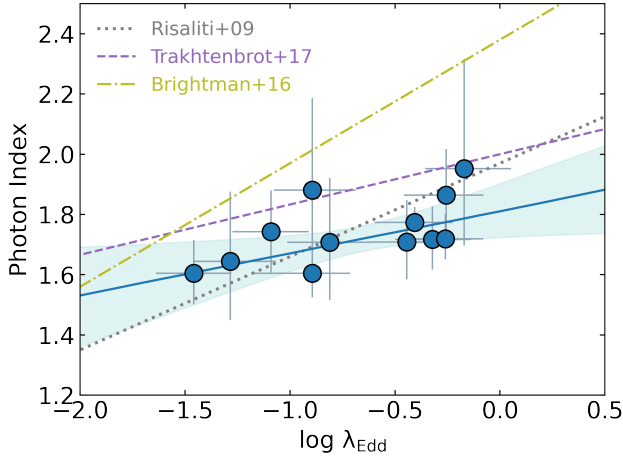


Figure 2. Eddington ratio (λ_{Edd}) versus photon index (Γ) from our sample (blue points), with the best-fit linear relation ($y = 0.14x - 1.85$) shown as a blue line. The shaded region represents the 1σ uncertainty. For comparison, the best-fit relations from Risaliti et al. 2009 (grey dotted line), Trakhtenbrot et al. 2017 (purple dashed line), and Brightman et al. 2016 (olive dotted-dashed line) are shown. Note that Trakhtenbrot et al. 2017 used three methods to compute the best fit; we report results for the two methods that agree with each other (see Trakhtenbrot et al. 2017 for details). Comparisons with other works are discussed in the text and omitted here for clarity.

This approach ensures proper propagation of asymmetric measurement errors into the regression analysis. The best-fit slope and intercept are $0.14^{+0.13}_{-0.11}$ and $1.85^{+0.11}_{-0.09}$, respectively, where the uncertainties are reported at 1σ . The standard deviation is 0.04 dex. These results are consistent when fitting with `linmix` (Kelly 2007), a hierarchical Bayesian model that accounts for error bars in both x and y directions. Since `linmix` requires symmetric uncertainties, we approximated them by averaging the asymmetric errors, yielding 0.14 ± 0.11 and $1.81^{+0.09}_{-0.08}$, for slope and intercept, respectively. Due to the limited number of data points available, we did not attempt to separate our sample into sub- and super-Eddington regimes.

Our results are shown in Figure 2. Within the errors, we found agreement with the results of Trakhtenbrot et al. (2017), who analyzed the full BASS DR1 sample, suggesting a similar behavior for both obscured AGNs (our sample) and the broader AGN population. Similarly, Brightman et al. (2016) analyzed a sample of obscured AGNs and found consistency with the relation reported by Brightman et al. (2013) for unobscured sources. The slopes reported by Brightman et al. (2013, 2016) are higher but consistent with our results, while we observed a discrepancy in the normalizations. For unobscured AGNs, studies such as Risaliti et al. (2009), Liu

et al. (2021), and Huang et al. (2020) reported steeper slopes that remain consistent with our results, albeit with slightly higher normalizations. Interestingly, our findings also align with the best-fit normalization reported by Kamraj et al. (2022), who used *NuSTAR* data but did not find a significant correlation. Nonetheless, the scatter of data points, uncertainties in the best-fit relations, different methods for determining Γ and λ_{Edd} , and variations in sample selection hinder a direct comparison of the results, with a deeper analysis of these effects lying beyond the scope of this study.

5.3. The obscuration and Eddington ratio forbidden region

Figure 3 shows the Eddington ratio λ_{Edd} versus the derived absorption N_{H} . The grey shaded region marks the so-called “forbidden region” where AGNs are not expected to persist due to the effects of radiation pressure on dusty gas from the central AGN engine (e.g., Fabian et al. 2008; Ricci et al. 2017b). When the Eddington ratio exceeds the effective Eddington limit for a given column density, radiation pressure overcomes gravitational forces, expelling the obscuring material and creating a ‘blowout’ region where long-lived dusty clouds cannot survive (e.g., Fabian et al. 2009). AGNs found within this region are likely caught in a short-lived transitional phase, during which radiation pressure actively reduces the column density (e.g., Kammoun et al. 2020; Ricci et al. 2022a, 2023a). This phase may involve variability in both accretion rate and obscuration, and can be associated with outflows/winds (e.g., Kakkad et al. 2016; Musiimenta et al. 2023).

In our sample, 6 out of 12 ($50 \pm 13\%$) sources (BAT IDs 555, 787, 1204, 1248, 1296, and 1586) are clearly in this region, with one additional source (BAT ID 32) lying very close to the threshold. This suggests that high-luminosity, obscured AGNs are frequently observed in this transitional phase. Among the sources in the forbidden region, two objects (BAT IDs 787 and 1248) show an outflow component in $[\text{O III}]\lambda 5007$ (Oh et al. 2022), with BAT ID 787 also showing winds in the X-rays (Tombesi et al. 2014; Mestici et al. 2024). Additionally, X-ray flux variability driven by changes in N_{H} is observed in BAT IDs 787, 1204, and 1296 (see Section 5.6), with two sources lacking multi-epoch data. Among the sources outside the forbidden region, outflows in $[\text{O III}]\lambda 5007$ are found in two sources (BAT IDs 80 and 199), and none show flux changes due to N_{H} variability, with the exception of one source for which we do not have multi-epoch data available.

While $[\text{O III}]\lambda 5007$ outflows are observed both inside and outside the forbidden region (two sources in each

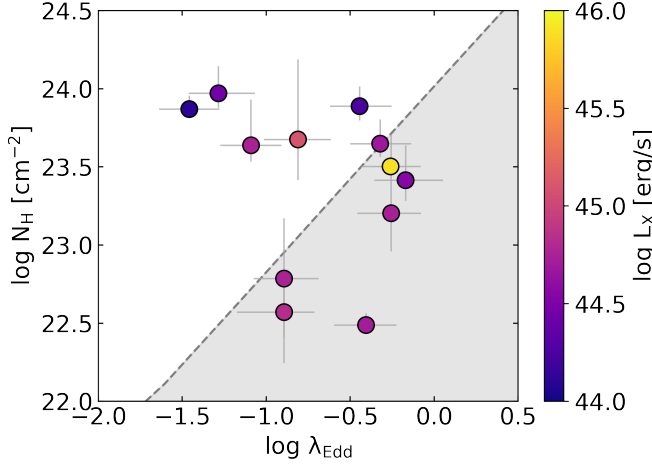


Figure 3. Eddington ratio (λ_{Edd}) versus absorption (N_{H}) for the 12 sources for which λ_{Edd} was computed, color-coded by X-ray luminosity (L_{X}). 6/12 sources are inside the so-called “forbidden” region (shaded grey), where AGNs are expected to be caught in a transitional phase where outflows/winds and/or changes in N_{H} are expected.

case), it is noteworthy that one of the outflow sources outside this area (BAT ID 199) lies not far away from the boundary, suggesting a potential connection between outflows and both the region itself and its immediate vicinity. Interestingly, another source near the boundary (BAT ID 32) has signs of activity, showing an ultrafast inflow in the X-rays (Peca et al. 2025). On the other hand, N_{H} variability is more prevalent for sources in the forbidden region. Indeed, among those with multi-epoch data, all but one within the region show variability, whereas none outside the region do. To summarize, the concentration of sources with N_{H} variability and the presence of inflows and outflows within or near the forbidden region support the interpretation that these AGNs are undergoing a transitional phase in their evolution, marked by increased and variable activity. However, we acknowledge that the relatively small number of objects in our sample prevents us from drawing more detailed statistical conclusions about these phenomena. Additional investigation of, for example, X-ray winds and outflows would be worthwhile to understand their role in the transitional phase further, but such analysis is beyond the scope of this paper.

These findings can be placed in the broader context of AGN obscuration models. Ananna et al. (2022b) presented a scenario where the simple geometric unified model applies to AGNs at low- λ_{Edd} regime, but at a transitional value of $\log \lambda_{\text{Edd}} \gtrsim -1.7$ (Fabian et al. 2008), the ratio of obscured to all AGNs reflects temporal evolution rather than just geometric evolution, i.e., it reflects how much time AGNs spend in the ob-

scured state. According to this model, at our median $\log \lambda_{\text{Edd}} = -0.59$, this ratio is about 20% for AGNs with $\log N_{\text{H}}/\text{cm}^{-2} \geq 22$. As the median $\log N_{\text{H}}$ of our sample is higher ($\simeq 23.5$), this may indicate the persistence of higher $\log N_{\text{H}}$ obscuring matter for a given λ_{Edd} . This is also in agreement with the Fabian et al. (2008) and Ricci et al. (2017c) scenario where the Compton-thick level of obscuration is expected to be persistent even at very high Eddington ratios ($\lambda_{\text{Edd}} \simeq 1$). Finally, we report no trend with luminosity, although we note that the three sources with $\log L_{\text{X}} < 44.5$ erg/s lie well outside this transitional region.

5.3.1. Comparison with Bär et al. 2019

Our work can be compared to Bär et al. (2019), who selected a sample of 28 luminous ($\log L_{\text{bol}}/\text{erg s}^{-1} \gtrsim 45.3$) and optical Seyfert 2 AGNs from the BASS DR1 survey. Six sources in our sample overlap with theirs. The differences between the samples arise from various factors, including the exclusion of 22 out of 28 sources from Bär et al. (2019) in our study due to their X-ray 2–10 keV luminosity falling below our chosen threshold. On the other hand, our sample includes 6 BASS DR1 sources not present in Bär et al. (2019), either because they were not yet classified in BASS DR1, were reclassified in BASS DR2, or were not included in the X-ray analysis of Ricci et al. (2017a).

Bär et al. (2019) showed that high-luminosity AGNs in general avoid the forbidden region, with a few exceptions. This is not in disagreement with our findings, as among the sources in our sample that lie in this region, only two overlap with the Bär et al. (2019) sample. These two sources (BAT IDs 555 and 1204) are located near the demarcation lines in our work, and the minor differences with Bär et al. (2019) likely arise from different fitting methods, datasets, and variability. Comparing the two full samples, we find consistent results with Bär et al. (2019), confirming that highly X-ray luminous and obscured AGN follow standard optical diagnostic diagrams and infrared classification schemes, as we further discuss in Appendix C.

5.4. Covering factor

The covering factor plays a crucial role in understanding the modeling of the AGN structure. Recent studies (e.g., Ricci et al. 2017c; Ananna et al. 2022a,b; Ricci et al. 2022a, 2023b) show that the covering factor varies significantly with the Eddington ratio, reflecting the dynamics between accretion processes and obscuring material. At low λ_{Edd} , AGNs tend to have larger covering factors, indicative of more extensive obscuration which may arise from dense, gravitationally bound clouds. As λ_{Edd}

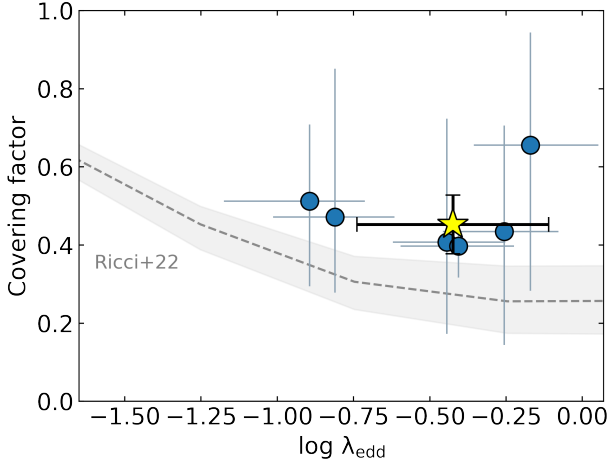


Figure 4. Eddington ratio (λ_{Edd}) versus covering factor for the sources where both measurements could be obtained (six sources). The yellow star represents the median value, and its uncertainties are computed as median absolute deviations. The grey dashed line is the relation by (Ricci et al. 2022a) with associated 1σ uncertainties.

increases, radiation pressure from the central source begins to push and expel the obscuring material, reducing the covering factor at higher accretion rates (e.g., Ricci et al. 2022a; Ananna et al. 2022b).

In this work, not all torus models allow for an estimate of the covering factor. Specifically, the MYTorus model does not provide such an estimate, as it assumes a fixed geometry for the torus, as described in Section 4.1. Due to this limitation, we could not estimate the covering factor for the eight sources where this model was preferred. For the remaining 13 sources, we obtained covering factor estimates using model-specific approaches. For the four sources modeled with RXTorusD, the covering factor was taken directly from the r/R ratio, a model parameter that directly represents the covering fraction (Ricci et al. 2023a). In UXCLUMPY, instead, the covering factor is not an explicit model parameter. However, it can be derived by interpolating the tabulated values provided by Boorman et al. (2024), which relate the $\text{TOR}\sigma$ and CTKcover parameters to the covering factor. For the nine sources in our sample modeled with this model, we applied this method and used the MCMC-derived posterior distributions of these parameters to estimate the associated uncertainties.

The median covering factor for the 13 sources where this measurement could be computed is 0.53 ± 0.09 . The uncertainty is estimated as the median absolute deviation, excluding one source for which only a lower limit was available. In Figure 4, we show our results against λ_{Edd} . This figure includes only six points, as λ_{Edd} was available only for these sources. We compare these re-

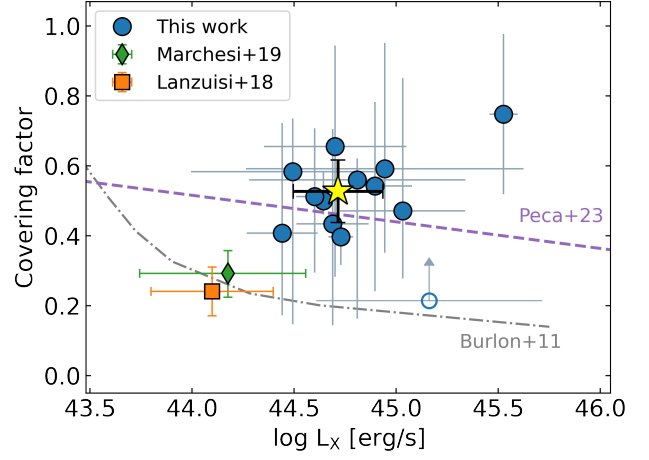


Figure 5. Intrinsic 2–10 keV X-ray luminosity (L_X) versus covering factor for the 13 sources in our sample where the latter could be estimated. The yellow star indicates the median value, with uncertainties are computed as median absolute deviations. The single lower limit (blue empty point) was excluded from the median calculation. For comparison, we also show the covering factor result from Marchesi et al. (2019) (green diamond), obscured Compton-thin fraction by Peca et al. (2023) (purple dashed line), and Compton-thick fractions by Lanzuisi et al. (2018) (orange square) and Burlon et al. (2011) (grey dotted-dashed line).

sults with the relation revised by Ricci et al. (2022a). Interestingly, our median value (yellow star) lies above the curve. This effect can be attributed to the fact that optically obscured AGNs tend to have higher covering factors than the broader AGN population (Ricci et al. 2011; Elitzur 2012; Ichikawa et al. 2015). However, we note that within the uncertainties all points are consistent with the expected relation.

Another well-studied relation is the one between the covering factor and X-ray luminosity (e.g., Brightman et al. 2015; Marchesi et al. 2019). The covering factor, being a direct geometrical proxy for the fraction of obscured AGNs, can be used to investigate the observed trend in which the obscured AGN fraction decreases with increasing luminosity (Treister & Urry 2005; Hasinger 2008; Aird et al. 2015; Buchner et al. 2015; Peca et al. 2023). This trend has been attributed to the combination of different mechanisms, including increasing radiation pressure, dust sublimation, geometrical effects, and the clearing of gas clouds (e.g., Hönig & Beckert 2007; Akylas & Georgantopoulos 2008; Hasinger 2008; Mateos et al. 2017; Matt & Iwasawa 2019).

In Figure 5 we show the derived covering factor versus X-ray luminosity. Notably, we extend earlier analyses (e.g., Marchesi et al. 2019; Zhao et al. 2021; Tanimoto et al. 2020, 2022) to higher intrinsic luminosities ($\log L_X/\text{erg s}^{-1} > 44.5$). For example, March-

esi et al. (2019) studied a sample of local Compton-thick AGNs at lower luminosities ($\log L_X/\text{erg s}^{-1} \sim 44$) and found lower covering factors compared to ours. When compared with other works on the fraction of obscured AGNs, our results show higher values than the Compton-thick fractions found by Lanzuisi et al. (2018) at $z < 1$ and Burlon et al. (2011) at $z < 0.3$, while aligning well with the Compton-thin ($22 < \log N_H/\text{cm}^{-2} < 24$) fractions from Peca et al. (2023) at $z < 1$ and other works as well (e.g., Aird et al. 2015; Buchner et al. 2015; Ananna et al. 2019). This agrees with the nature of our sample, which contains heavily obscured AGNs, though only one source exceeds the Compton-thick threshold. These comparisons support the scenario in which the covering factor of the Compton-thin material is larger than that of the Compton-thick gas (e.g., Tanimoto et al. 2020; Zhao et al. 2021). Finally, we do not report any trend with luminosity, likely due to the limited luminosity range of our sample.

5.5. Iron $K\alpha$ Emission Lines

We proceeded with the characterization of the iron $K\alpha$ emission lines as follows. The torus models selected for the spectral analysis include the iron line in a self-consistent manner, meaning it is built into the model and not treated as an independent, free parameter. To address this limitation, we performed a new fit using the MYTorus model in its simpler, coupled configuration, disabling the emission lines module. A Gaussian line at a fixed rest-frame energy of 6.4 keV was then manually added. We assumed a fixed width of $\sigma=10$ eV, consistent with an unresolved component (e.g., Vignali et al. 2015; Nanni et al. 2018). We verified that the fitted power-law continuum in this configuration was consistent, within uncertainties, with the best-fit model preferred by the DIC criterion for the main analysis. This ensured that the measurements of the emission line were performed on a model consistent with the one used for the spectral analysis. We considered the emission line significant if the difference in the fit statistic between the model with and without the Gaussian line was $\Delta\chi^2 > 2.7$, corresponding to a 90% confidence level for the additional line normalization free parameter. We computed the rest-frame equivalent width for the 9/19 ($47 \pm 11\%$) sources where this criterion was met. For the remaining sources, we computed the 3σ upper limits. We do not report any values for the two sources located in galaxy clusters (BAT ID 1204 and 1346), as their spectra are dominated by thermal emission from the *apex* component, which includes multiple emission lines, thereby making the measurement of the iron $K\alpha$ line unreliable. The results are shown in Table 3. Note

that for GALEXMSC J025952.92+245410.8 (BAT ID 1248), the values reported for the iron $K\alpha$ line are based solely on the *NuSTAR* spectra, as the *Chandra* spectrum has a limited number of counts, causing the line significance to fall below the detection threshold.

BAT ID	$K\alpha$ EW (eV)	$\Delta\chi^2$
20	< 0.18	< 2.7
32	$0.10^{+0.16}_{-0.05}$	22.2
80	$0.28^{+0.17}_{-0.10}$	24.3
119	< 0.19	< 2.7
199	$0.11^{+0.07}_{-0.07}$	6.8
476	$0.08^{+0.05}_{-0.04}$	15.1
494	< 0.46	< 2.7
505	$0.32^{+0.05}_{-0.24}$	15.5
555	< 0.17	< 2.7
714	$0.19^{+0.09}_{-0.05}$	22.5
787	< 0.04	< 2.7
1204	...	...
1241	< 0.56	< 2.7
1248	$0.18^{+0.18}_{-0.16}$	4.0
1291	< 0.89	< 2.7
1296	$0.10^{+0.03}_{-0.06}$	9.0
1346	...	...
1515	$0.10^{+0.09}_{-0.05}$	11.8
1586	< 0.23	< 2.7
1595	< 0.24	< 2.7
1630	< 1.08	< 2.7

Table 3. Rest-frame Fe $K\alpha$ equivalent width (EW) measurements using the MYTorus model (see text for details). For sources where the inclusion of the emission line led to a significant fit improvement at the 90% confidence level ($\Delta\chi^2 > 2.7$) we report the measured EW, while for the remaining sources we provide 3σ upper limits. No values are reported for BAT ID 1204 and 1346, as their spectra are dominated by thermal emission, preventing any line measurement. Columns, from left to right: source name, rest-frame EW, and $\Delta\chi^2$ between models with and without the emission line.

Figure 6 (left panel) shows that the EWs from our work are consistent with predictions made by different models when compared against N_H . Specifically, we compare our measurements with the e-torus model from Ikeda et al. (2009) and with the MYTorus model Murphy & Yaqoob (2009b), considering all possible inclination angles. Additionally, we compare our results to the empirical best-fit relation derived by Guainazzi et al. (2005) from a sample of Seyfert 2 galaxies at $z < 0.1$. Our data confirm their findings, showing consistency with a correlation between EW and N_H for column den-

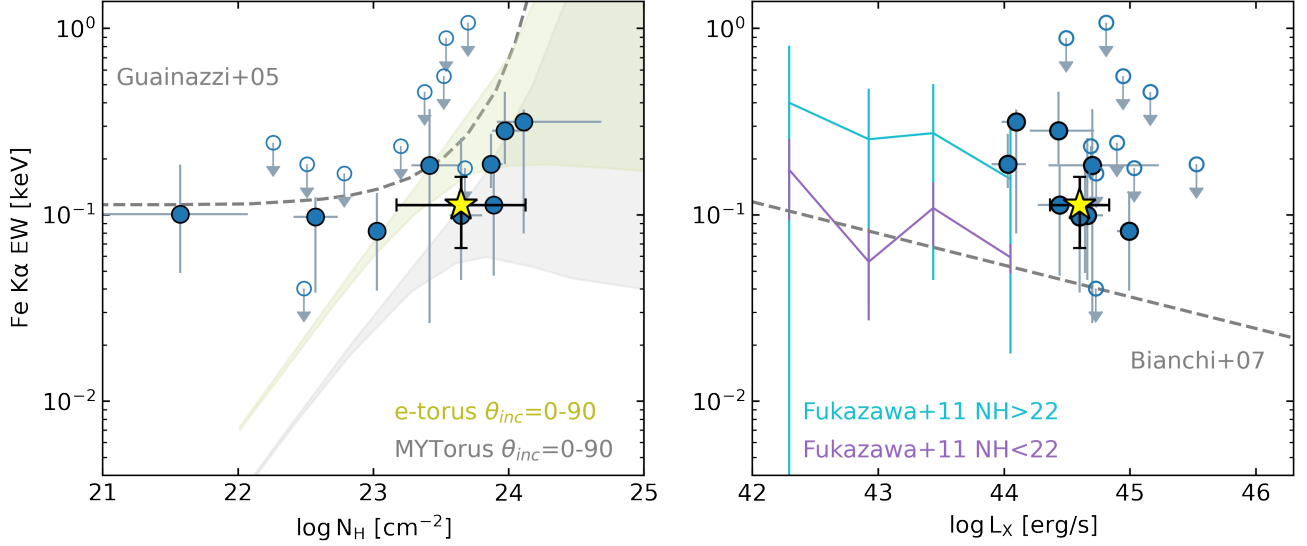


Figure 6. *Left:* Fe K α rest-frame equivalent width as a function of the derived N_{H} . The shaded areas represent model predictions from e-torus (Ikeda et al. 2009) (olive) and MYTorus (Murphy & Yaqoob 2009b) (grey), showing the range of all possible inclination angles. The torus opening angles assumed in the models are 30° and 60°, respectively. The empirical relation from Guainazzi et al. (2005) is shown as a grey dashed line. *Right:* Fe K α rest-frame equivalent width as a function of X-ray luminosity. Cyan and purple points show binned results from Fukazawa et al. (2011) for obscured and unobscured AGNs, respectively. Each point represents the average in a given L_{X} bin, with error bars corresponding to one standard deviation. The dashed grey line indicates the relation found by Bianchi et al. (2007) for unobscured AGNs. In both panels, empty points indicate 3σ upper limits, where error bars on the x-axis are omitted for clarity but correspond to ~ 0.5 dex. No values are reported for BAT ID 1204 and 1346, as their spectra are dominated by thermal emission from the *apec*, preventing any measurements of the emission line. The yellow stars represent the median values of the constrained values, with uncertainties computed as median absolute deviations.

sities above $\log N_{\text{H}}/\text{cm}^{-2} \sim 23$, and a flattening at lower values around EWs of about 100 eV. This flattening is also consistent with the results of Fukazawa et al. (2011), who measured iron K α EWs within the range of [50–120] eV for sources with $\log N_{\text{H}}/\text{cm}^{-2} < 23$. Such behavior supports an origin of the iron line emission from the inner walls of the torus or the broad line region (e.g., Ghisellini et al. 1994; Guainazzi et al. 2005; Fukazawa et al. 2011). Our median EW lies slightly below the Guainazzi et al. (2005) relation, which can be explained by the lower luminosity of their sample ($\log L_{\text{X}}/\text{erg s}^{-1} \sim 42$), as discussed in what follows.

In Figure 6 (right panel), we compare our EW results against L_{X} . Many established relations in the literature connect X-ray luminosity with the equivalent width of the iron K α line, for both unobscured (e.g., Iwasawa & Taniguchi 1993; Bianchi et al. 2007) and obscured (Ricci et al. 2014b,a; Boorman et al. 2018) AGNs. This trend, known as the X-ray Baldwin Effect (Iwasawa & Taniguchi 1993), takes its name from the discovery of its analog in the UV for the CIV λ 1549 line (Baldwin 1977). In X-rays, this effect is attributed to factors such as the ionization of circumnuclear material, which weakens line emission, continuum dilution with increasing luminosity, and a luminosity-dependent decrease in

the torus covering factor and line-of-sight N_{H} , which diminishes reprocessed emission (e.g., Ricci et al. 2014b,a; Boorman et al. 2018). We compare our measurements with the relation from Bianchi et al. (2007), which was derived for unobscured AGNs, and the values reported by Fukazawa et al. (2011), who analyzed both obscured and unobscured AGNs using *Suzaku* data. Although the limited number of data points prevents us from looking for possible correlations, our results are more consistent with those of obscured AGNs from Fukazawa et al. (2011). This aligns with expectations, as obscured AGNs typically exhibit higher equivalent widths due to the suppression of the continuum by absorption and the stronger contribution of reprocessed emission from dense circumnuclear material (e.g., Ricci et al. 2014b,a; Lanzuisi et al. 2015a; Peca et al. 2021). As discussed in Ricci et al. (2014b), this enhancement in EW for obscured AGNs is primarily an observational effect, potentially biasing its use as a diagnostic in studies of the Baldwin Effect.

5.6. Variability analysis

In the previous sections, we characterized our sample using the average values of the spectral parameters L_{X} and N_{H} . Among the 21 sources, 13 have multi-epoch

BAT ID	Flux (erg/s/cm ²)	C_{AGN}	Log N_H (cm ⁻²)	Date	Instr.	p-value (Flux/ N_H)
32	18.9 ^{+0.5} _{-0.9}	1	23.7 ^{+0.2} _{-0.1}	2010-12-19	Suzaku	2e-13/0.93
	15.6 ^{+0.9} _{-1.0}	0.8 ^{+0.8} _{-0.7}	23.6 ^{+0.2} _{-0.2}	2023-05-08	NuSTAR	
	9.6 ^{+0.9} _{-0.7}	0.6 ^{+0.7} _{-0.5}	23.7 ^{+0.2} _{-0.2}	2024-08-22	NuSTAR	
80	5.3 ^{+0.3} _{-0.3}	1	23.9 ^{+0.3} _{-0.2}	2016-12-03	NuSTAR	3e-10/0.95
	3.6 ^{+0.2} _{-0.5}	1.1 ^{+0.4} _{-0.3}	24.1 ^{+0.3} _{-0.2}	2017-10-31	XMM	
	4.9 ^{+0.2} _{-0.6}	1.0 ^{+0.4} _{-0.3}	23.9 ^{+0.3} _{-0.2}	2020-08-30	NuSTAR	
	5.9 ^{+0.4} _{-0.9}	1.5 ^{+0.6} _{-0.4}	24.0 ^{+0.3} _{-0.2}	2021-06-17	XMM	
119	86.0 ^{+3.5} _{-3.3}	1	22.5 ^{+0.2} _{-0.1}	2006-01-23	XMM	<1e-16/0.95
	27.9 ^{+5.5} _{-7.9}	0.3 ^{+0.2} _{-0.1}	22.5 ^{+0.2} _{-0.4}	2024-12-18	NuSTAR	
199	10.3 ^{+0.5} _{-0.6}	1	23.9 ^{+0.2} _{-0.2}	2016-05-06	NuSTAR	3e-7/0.99
	6.5 ^{+0.4} _{-0.7}	0.6 ^{+0.5} _{-0.3}	23.9 ^{+0.2} _{-0.2}	2017-11-20	XMM	
	8.6 ^{+0.5} _{-0.5}	0.9 ^{+0.4} _{-0.3}	23.9 ^{+0.2} _{-0.2}	2021-06-04	NuSTAR	
476	38.0 ^{+0.6} _{-1.0}	1	23.1 ^{+0.1} _{-0.1}	2009-01-20	XMM	2e-10/4e-2
	44.9 ^{+1.5} _{-1.9}	1.1 ^{+0.1} _{-0.1}	22.9 ^{+0.1} _{-0.2}	2023-07-12	NuSTAR	
505	3.7 ^{+0.3} _{-0.5}	1	23.9 ^{+1.0} _{-0.2}	2018-06-02	NuSTAR	0.43/0.82
	3.2 ^{+0.2} _{-0.6}	0.8 ^{+1.0} _{-0.4}	24.0 ^{+0.5} _{-0.1}	2018-06-02	XMM	
	3.4 ^{+0.7} _{-0.2}	1.3 ^{+0.9} _{-0.5}	24.3 ^{+0.6} _{-0.5}	2020-06-24	NuSTAR	
714	7.0 ^{+0.7} _{-0.7}	1	23.9 ^{+0.2} _{-0.1}	2016-05-25	NuSTAR	4e-7/0.80
	11.8 ^{+0.4} _{-1.1}	1.4 ^{+0.3} _{-0.2}	23.9 ^{+0.1} _{-0.1}	2018-02-09	XMM	
	12.0 ^{+0.5} _{-0.6}	1.3 ^{+0.3} _{-0.3}	23.8 ^{+0.1} _{-0.1}	2020-03-25	NuSTAR	
787	57.0 ^{+45.5} _{-45.9}	1	22.7 ^{+0.1} _{-0.1}	2008-09-21	XMM	0.76/4e-13
	74.0 ^{+10.9} _{-34.8}	1.0 ^{+0.1} _{-0.2}	22.2 ^{+0.3} _{-0.2}	2016-07-12	NuSTAR	
1204	54.8 ^{+0.6} _{-0.1}	1	23.6 ^{+0.1} _{-0.1}	2013-11-13	XMM	<1e-16/3e-2
	56.2 ^{+0.2} _{-0.5}	1.1 ^{+0.1} _{-0.1}	23.6 ^{+0.1} _{-0.1}	2013-11-22	XMM	
	64.0 ^{+1.2} _{-1.9}	1.1 ^{+0.1} _{-0.1}	23.3 ^{+0.2} _{-0.2}	2021-01-07	NuSTAR	
1248	10.7 ^{+1.2} _{-4.0}	1	23.5 ^{+0.3} _{-0.2}	2023-03-26	NuSTAR	0.82/0.68
	11.2 ^{+8.0} _{-9.8}	0.8 ^{+0.5} _{-0.3}	23.3 ^{+0.3} _{-0.2}	2024-11-26	Chandra	
1296	28.4 ^{+0.8} _{-2.0}	1	21.8 ^{+0.4} _{-0.5}	2023-02-12	NuSTAR	<1e-16/7e-7
	14.8 ^{+0.4} _{-0.7}	0.7 ^{+0.8} _{-0.6}	22.8 ^{+0.1} _{-0.1}	2025-03-14	XMM	
1346	22.8 ^{+0.7} _{-0.5}	1	23.7 ^{+0.4} _{-0.1}	2007-04-13	XMM	3e-4/0.57
	26.6 ^{+1.0} _{-0.8}	1.1 ^{+0.2} _{-0.1}	23.6 ^{+0.4} _{-0.2}	2024-11-25	NuSTAR	
1515	22.2 ^{+1.9} _{-5.1}	1	21.9 ^{+0.5} _{-1.0}	2023-03-11	NuSTAR	4e-9/ -
	32.1 ^{+2.6} _{-1.0}	1.4 ^{+0.4} _{-0.3}	< 21.1	2025-02-27	XMM	

Table 4. Summary of parameters left free to vary in simultaneous fits for sources with multi-epoch data. From the *left*: BAT ID, observed 2–10 keV best-fit flux in units of 10^{-13} erg/s/cm², C_{AGN} (fixed to 1 for the first observation of each source), line-of-sight column density N_H , observation date, telescope, and p-values on flux and N_H obtained from Equation 9. The value <1e-16 indicates a p-value lower than the limits from our Python scripts. For BAT ID 1515, we computed the p-values only for the flux, as N_H is constrained only as an upper limit in the second observation (see text for details).

spectroscopy: 11 observed with XMM-*Newton* and NuSTAR, one with Chandra and NuSTAR, and one observed twice with NuSTAR and one with Suzaku. Here, we focus on these 13 sources and present single-epoch results. Specifically, we examine the two primary parameters left free to vary in the spectral fitting: C_{AGN} , which accounts for intrinsic flux variability, and the line-of-sight N_H . We also include the observed fluxes in the 2–10 keV band. The temporal evolution of these parameters is

summarized in Table 4, with a visual characterization shown in Figures 19, 20, and 21 in Appendix A.

While a detailed study of the variability of these parameters is beyond the scope of this paper and it is the subject of ongoing efforts, we quantified the variability using the method proposed by Torres-Albà et al. (2023). This approach, which has been used to evaluate variability in N_H and flux (Torres-Albà et al. 2023; Sengupta et al. 2025; Pizzetti et al. 2025), employs a chi-squared-

like procedure:

$$\chi^2 = \sum_{i=1}^n \frac{(N_{H,\text{los},i} - \langle N_{H,\text{los}} \rangle)^2}{\delta(N_{H,\text{los},i})^2} \quad (9)$$

where the “true” value is the average value from the available i -observations, and the uncertainty δ is defined as the lower uncertainty when the i -value is less than the average or the upper uncertainty otherwise. The p-value is then computed by assuming chi-squared statistics. A source is considered variable if the p-value is less than 0.05 (Torres-Albà et al. 2023; Pizzetti et al. 2025). We applied this method to the observed flux and N_H . We did not perform this computation for C_{AGN} , since its value is fixed to unity for the first observation in the simultaneous fitting procedure, limiting the method’s applicability. Our results are presented in Table 4, where we also report the derived C_{AGN} values, as they still provide qualitative support for possible intrinsic variability. For one source only (BAT ID 1515), we applied the method only on the flux, as the N_H value for the second observation is an upper limit. This is due to a combination of the source being unobscured ($\log N_H/\text{cm}^{-2} < 22$) and lying behind a high Galactic column density ($N_H = 1.2 \cdot 10^{21} \text{ cm}^{-2}$, Kalberla et al. 2005), preventing our ability to constrain N_H .

Overall, 11/13 sources ($85^{+5}_{-15}\%$) exhibit variability in either flux (10/13, $77^{+8}_{-15}\%$) or N_H (4/12, $33^{+15}_{-10}\%$). Additionally, 3/12 ($25^{+16}_{-8}\%$) are variable in both flux and N_H , 6/12 ($50^{+13}_{-13}\%$) show flux variability without changes in N_H , and 1/12 ($8^{+15}_{-3}\%$) displays variability in N_H only. These distinctions suggest that AGN variability may be driven by changes in N_H , intrinsic flux, or a combination of both processes. It is worth mentioning that our results on the fraction of sources with N_H variability are in excellent agreement to what was found by Torres-Albà et al. (2023) and Pizzetti et al. (2025) in a sample of lower luminosity ($\log L_X/\text{erg s}^{-1} \sim 43$) local Compton-thin AGNs (Zhao et al. 2021).

6. RESULTS II: [Ne v] $\lambda 3426$ EMISSION LINE

6.1. [Ne v] $\lambda 3426$ and X-ray obscuration

The [Ne v] $\lambda 3426$ emission line is a forbidden transition with a high-ionization potential of $h\nu > 97 \text{ eV}$. As such energetic photons are rarely produced by stellar processes, [Ne v] $\lambda 3426$ is considered a reliable tracer of AGN activity (e.g., Zakamska et al. 2003; Gilli et al. 2010; Yuan et al. 2016; Negus et al. 2023; Cleri et al. 2023a). Unlike X-ray emission, which is produced near the central SMBH and may be heavily attenuated by the obscuring torus, the [Ne v] $\lambda 3426$ line originates in the narrow-line region, located at larger scales (e.g., Hickox & Alexander 2018; Negus et al. 2023). As a

result, [Ne v] $\lambda 3426$ emission can remain observable even in heavily obscured AGNs where X-rays are suppressed, making it a valuable tracer for uncovering such systems (e.g., Mignoli et al. 2013; Vergani et al. 2018; Mignoli et al. 2019). In this context, combining X-ray and [Ne v] $\lambda 3426$ emissions offers a powerful tool for probing obscuration, as the X-ray-to-[Ne v] $\lambda 3426$ flux ratio can be used as a direct indicator of the absorbing column density (e.g., Gilli et al. 2010; Vignali et al. 2014; Lanzuisi et al. 2015b; Barchiesi et al. 2024; Cavicchi et al. 2025). Although powerful, it is worth mentioning that this method can be influenced by dust, gas attenuation, extreme star formation in the host galaxy, and geometrical effects, potentially affecting both emission components (Buchner & Bauer 2017; Gilli et al. 2022; Cleri et al. 2023a, e.g.). With this in mind, in this section we investigate the [Ne v] $\lambda 3426$ emission line in our sample of highly luminous, obscured AGNs and explore its relation with the X-ray emission.

For this study, we obtained emission line flux measurements from a compilation of optical observations performed with the VLT/X-shooter and KCWI instruments, as described in Section 2.3. Optical data were available for all but one source (BAT ID 119), which was observed under poor conditions, making it unsuitable for detailed analysis. The uncertainties on line measurements are less than 0.1 dex (Oh et al. 2022; Koss et al. 2022b), with overall uncertainties primarily driven by systematic calibration errors. To account for this, we added a 20% systematic calibration error to the uncertainties provided by the pipeline (see Section 2.3) on the line measurements. We detected [Ne v] $\lambda 3426$ emission in 17/20 sources, corresponding to a detection rate of $85^{+5}_{-11}\%$. For the three additional sources, we derived 3σ upper limits. Our detection rate is slightly higher but consistent with Reiss et al. (2025), hereafter R25, which analyzed narrow-line AGNs in the full BASS DR2 sample to study the [Ne v] $\lambda 3426$ emission and found a detection rate as high as $\sim 70\%$ for high SNR spectra.

Our results can be compared with previous studies at slightly higher redshifts. For instance, Barchiesi et al. (2024), analyzing a sample of [Ne v] $\lambda 3426$ emitters at $0.6 < z < 1.2$, found that two-thirds are heavily X-ray obscured ($\log N_H/\text{cm}^{-2} > 23$) and that half are undergoing intense accretion ($\log \lambda_{\text{Edd}} > -1$). These findings are consistent with our results, as we detect a high fraction of [Ne v] $\lambda 3426$ emission among sources that are predominantly obscured and rapidly accreting, with median values of $\log N_H/\text{cm}^{-2} \simeq 23.5$ and $\log \lambda_{\text{Edd}} \simeq -0.6$, respectively. This agreement is particularly noteworthy given that half of our sample lies within the forbidden region of the N_H - λ_{Edd} parameter space, where

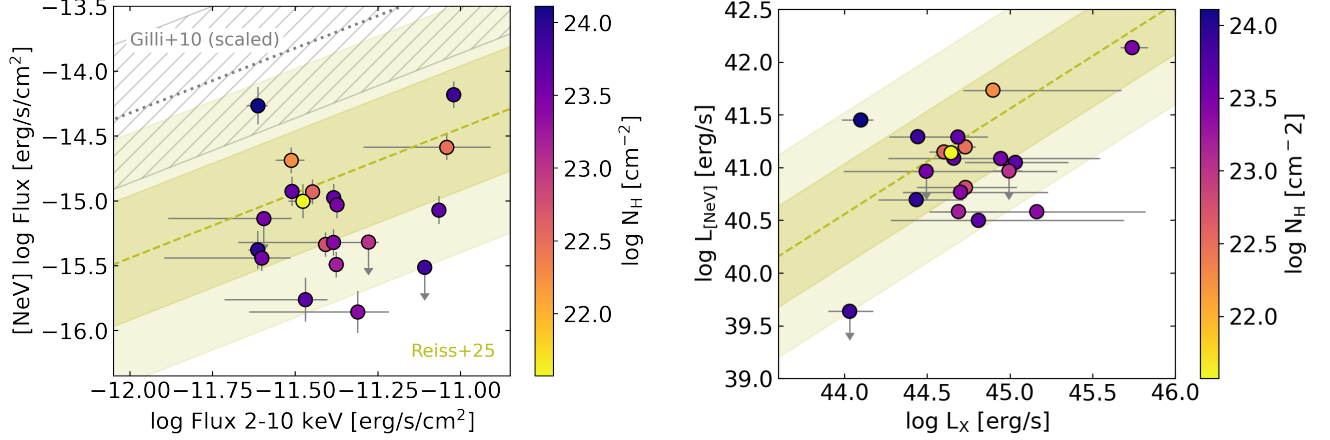


Figure 7. *Left panel:* $[\text{Ne v}] \lambda 3426$ flux versus absorption-corrected 2–10 keV flux for our sample. For comparison, we show the R25 median relation for the BASS DR2 sample (olive dashed line), along with its 1σ and 2σ uncertainties (olive shaded areas), and the Compton-thick AGN threshold from Gilli et al. (2010) with its 1σ uncertainty (grey dotted line and hatched area). Since Gilli et al. (2010) used observed fluxes (i.e., uncorrected for absorption), we applied a correction factor of 14 (see details in the text) to align with our intrinsic fluxes. *Right panel:* Intrinsic (i.e., de-absorbed 2–10 keV rest-frame) luminosity versus $[\text{Ne v}] \lambda 3426$ luminosity. We compare our results with the R25 median value relation with its 1σ and 2σ uncertainties (olive dashed line and shaded areas). In both panels, data points are color-coded by the line-of-sight column density (N_H). The $[\text{Ne v}] \lambda 3426$ flux is not corrected for extinction or reddening.

radiation pressure is expected to clear out the obscuring material (see Section 5.3), therefore demonstrating that $[\text{Ne v}] \lambda 3426$ can be detected in sources in this evolutionary phase. Furthermore, both Barchiesi et al. (2024) and Vergani et al. (2018), who studied another sample of $[\text{Ne v}] \lambda 3426$ emitters at similar redshifts, reported enhanced star formation rates compared to non- $[\text{Ne v}] \lambda 3426$ emitters and interpreted this as a signature of the star formation quenching phase. Our sample is consistent with these studies, as we find comparable stellar masses, with a median $\log M_*/M_\odot = 10.9$ based on the nine sources with available measurements (Koss et al. 2022a), and star formation rates, with a median $\sim 30 M_\odot \text{yr}^{-1}$ based on the seven sources with values reported by Ichikawa et al. (2017). Additionally, our $[\text{Ne v}] \lambda 3426$ luminosities are consistent with the $[\text{Ne v}] \lambda 3426$ luminosity–stellar mass relations found by Vergani et al. (2018).

Figure 7 presents the $[\text{Ne v}] \lambda 3426$ flux versus the observed 2–10 keV X-ray flux (left panel) and the de-absorbed, 2–10 keV rest-frame X-ray luminosity versus the $[\text{Ne v}] \lambda 3426$ luminosity (right panel). Both panels are color-coded by the N_H derived from our results presented in Section 5. For $[\text{Ne v}] \lambda 3426$ fluxes, we accounted for Galactic extinction following the prescriptions by Oh et al. (2022) but did not correct for intrinsic extinction, as our goal is to examine the observed properties of $[\text{Ne v}] \lambda 3426$ and ensure a meaningful comparison with literature results. We first compare our results with R25, which found $\log(L_{[\text{Ne v}]} / L_X) \simeq -3.36$ in

the BASS DR2 sample. Our luminosity plot shows less scatter than the flux plot, which is in line with their findings. Both distributions appear slightly skewed toward higher X-ray and lower $[\text{Ne v}] \lambda 3426$ flux and luminosity. However, all our points are consistent when considering the 2σ uncertainty on the R25 relation, indicating overall agreement. Furthermore, R25 have several upper limits for the $[\text{Ne v}] \lambda 3426$ flux range covered in this work, which could potentially influence their relation at these fluxes. We also confirm the absence of any correlation with N_H . Next, we compare our results with the Compton-thick AGN selection criterion from Gilli et al. (2010), which identifies Compton-thick candidates as those with an X-ray-to- $[\text{Ne v}] \lambda 3426$ flux ratio below 15. To align with our analysis, we rescaled their observed fluxes to intrinsic fluxes by applying a correction factor of 14, determined by assuming a power-law with $\Gamma = 1.8$ and $N_H = 10^{24} \text{cm}^{-2}$. Our sample appears skewed toward lower $[\text{Ne v}] \lambda 3426$ fluxes and higher X-ray fluxes, even when accounting for the fact that our median N_H is slightly lower than the Compton-thick level. A key difference is that our sample comprises higher-luminosity AGNs, particularly at low redshifts, resulting in higher fluxes for the same level of N_H . Moreover, the Gilli et al. (2010) sample may tend toward higher $[\text{Ne v}] \lambda 3426$ fluxes, as the authors pointed out. Consequently, we conclude that the apparent discrepancy is likely attributable to differences in sample properties and selection criteria. Despite these differences, we note that the only object in our sample that is nom-

inally Compton-thick (BAT ID 505) falls within the 1σ region of the [Gilli et al. \(2010\)](#) relation.

6.2. $[\text{Ne V}]\lambda 3426$ stacking and comparison with JWST

This section aims to connect our sample with the high-redshift ($z > 2$) Universe, particularly with JWST results. To this end, we used the emission-line best-fit spectra from [Oh et al. \(2022\)](#) for the sources available in BASS DR2, and derived the remaining using the same fitting procedure for consistency across all 21 sources in our sample. The template spectra were stacked using a bootstrap procedure with 1000 iterations. We then matched our stacked spectrum with the stacking analysis of [Mazzolari et al. \(2024\)](#), hereafter [M24](#), who studied 52 narrow-line AGNs in the CEERS survey ([Finkelstein et al. 2022](#)) at $2 \lesssim z \lesssim 9$. As in [M24](#), we renormalized our stack using the $[\text{O III}]\lambda 5007$ emission line peak flux. Because our spectra have higher resolution ($R > 1000$; see [Koss et al. 2022b](#)), compared to the lower resolution of the JWST stacked spectrum ($R \sim 1000$), we degraded our data to match the resolution of [M24](#) before stacking. To do this, we used the `SpectRes` code ([Carnall 2017](#)), which resamples spectra and their associated uncertainties onto an arbitrary wavelength grid. This function works with any grid of wavelength values, including non-uniform sampling, and preserves the integrated flux. The stacking results are shown in Figure 8. We note that, while the continuum was fitted and subtracted in [M24](#), here we used the emission line best-fit templates from [Oh et al. \(2022\)](#).

Interestingly, the only emission lines completely missing in the CEERS stacked spectrum are those from $[\text{Ne V}]$ doublet ($[\text{Ne V}]\lambda 3426$ and $[\text{Ne V}]\lambda 3346$). While other differences are visible in the residual spectrum (Figure 8 bottom panel), they primarily arise from variations in line widths or peak fluxes rather than the absence of specific emission lines. As shown in the zoom-in of Figure 8, our stack clearly exhibits a strong $[\text{Ne V}]\lambda 3426$ emission line, while the CEERS stack does not. Additionally, the $[\text{Ne V}]\lambda 3346$ emission line is absent. However, the lack of $[\text{Ne V}]\lambda 3346$ is not unexpected, as it is too faint to be detected and consistent with the noise level. In contrast, the $[\text{Ne V}]\lambda 3426$ emission line strength from stacking our sample is up to six times higher than the noise level in the JWST stacked spectrum, where the noise was estimated as $\text{rms} \sim 0.01$ over the 3250–3700 Å range, a region free of emission lines. These results are consistent when performing the stacking using the best spectroscopic resolution of the spectra in our sample.

This discrepancy becomes particularly intriguing when we consider the quantitative expectations based on

established scaling relations. The average expected ratio $\log(L_X/L_{[\text{O III}]})$ for type 2 AGNs is $\simeq 2.1$ (e.g., [Georgantopoulos & Akylas 2010](#); [Berney et al. 2015](#); [Ueda et al. 2015](#); [Malkan et al. 2017](#); [Oh et al. 2022](#)). When combined with the relation from [R25](#), $\log(L_{[\text{Ne V}]} / L_X) \simeq -3.36$, we obtain an expected $L_{[\text{Ne V}]} / L_{[\text{O III}]}$ ratio of ~ 0.06 . Such a ratio corresponds to a $[\text{Ne V}]\lambda 3426$ flux well above the JWST noise level and matches what we observe in our stacked spectrum. We note that these scaling relations span a wide range of bolometric luminosities, including values below those probed in our study and consistent with those found in [M24](#) and other JWST studies of obscured AGNs at high redshift (e.g., [Yang et al. 2023](#); [Lyu et al. 2024](#); [Maiolino et al. 2025](#)), typically around $\log L_{\text{bol}} / \text{erg s}^{-1} \approx 44\text{--}45$. This supports the expectation that the observed $[\text{O III}]\lambda 5007$ emission in JWST spectra should be accompanied by detectable $[\text{Ne V}]\lambda 3426$ emission, despite the relations being calibrated at low redshift and subject to possible evolutionary effects.

The presence of detectable $[\text{Ne V}]\lambda 3426$ emission is particularly relevant given the debate surrounding high-redshift AGNs identified by JWST, many of which exhibit very weak or entirely undetected X-ray emission (e.g., [Ananna et al. 2024](#); [Maiolino et al. 2025](#); [Mazzolari et al. 2024](#)). As we know from prior studies that X-ray obscuration increases with redshift (e.g., [Aird et al. 2015](#); [Buchner et al. 2015](#); [Vito et al. 2018](#); [Vijarnwanaluk et al. 2022](#); [Peca et al. 2023](#); [Signorini et al. 2023](#); [Pouliasis et al. 2024](#)), we can invoke high levels of obscuration to explain their observed X-ray weakness. In this context, the $[\text{Ne V}]\lambda 3426$ emission line can be exceptionally useful as an obscured AGN indicator, given its high detection rate in such systems, as shown both in this work and by [R25](#). We explore the implications of these findings in further detail in Section 7.

7. DISCUSSION

This work provides a comprehensive X-ray study of high luminosity ($L_X \gtrsim 10^{44}$ erg/s), optically obscured (Seyfert 1.9 and 2) AGNs at redshifts $z < 0.6$. In addition to detailed X-ray spectral modeling, we analyzed the UV/optical $[\text{Ne V}]\lambda 3426$ emission line associated with the X-ray sources, enabling a multi-wavelength perspective on this obscuration indicator. In this section, we discuss the broader implications of our findings.

7.1. The high luminosity and high obscuration regime

In Figure 9, we compare our sample with other X-ray surveys, showing only sources at $z < 0.6$ to match the redshift range of our sample. The broad region covered by our full sample is outlined with black dotted lines and corresponds to $\log L_X / \text{erg s}^{-1} > 44$ and

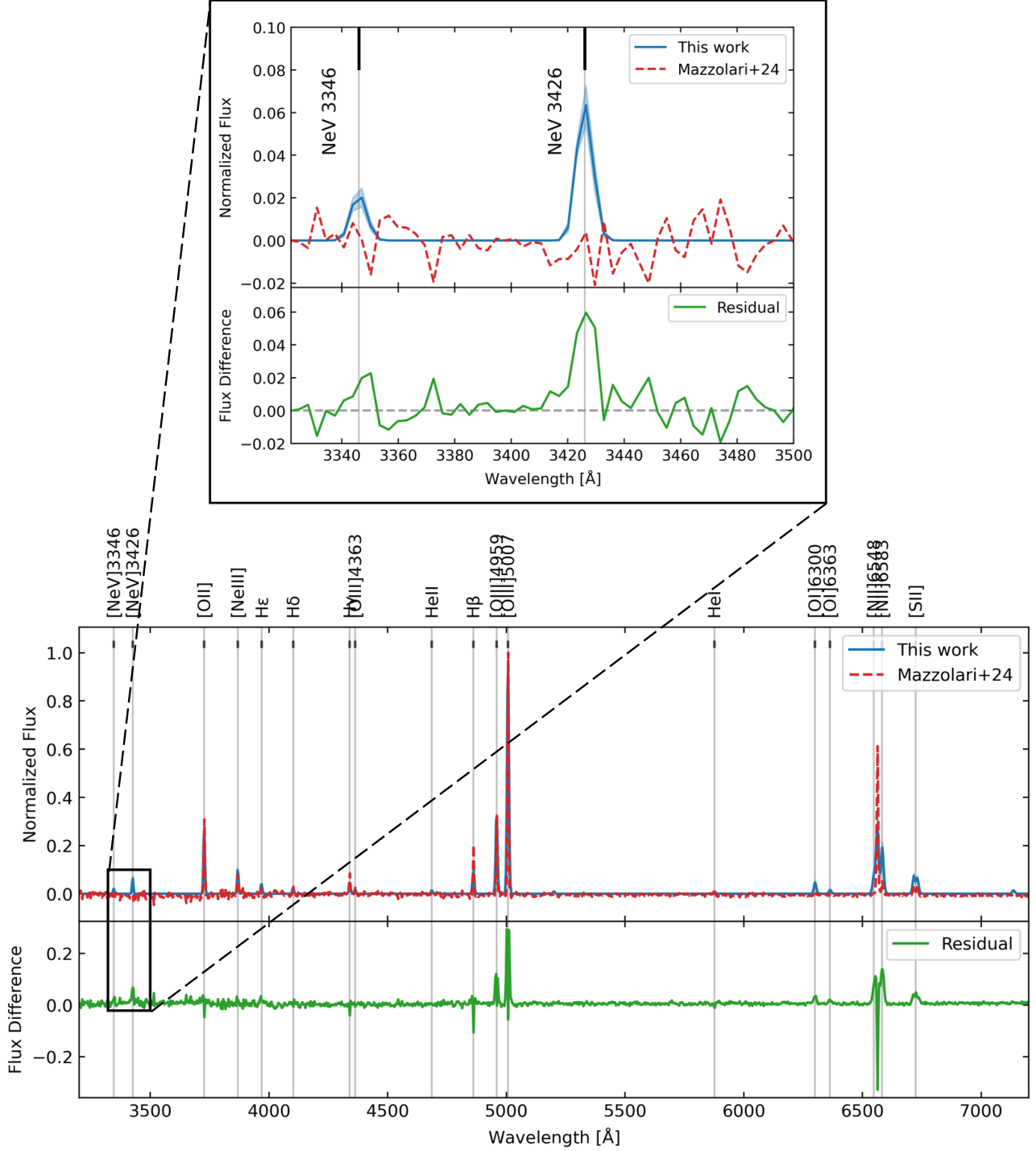


Figure 8. Comparison between the stacked high-redshift JWST spectrum from M24 (red dashed lines) and our stacked spectrum of low-redshift, highly luminous BASS AGNs (blue solid lines). The shaded blue region shows the 1σ uncertainty derived via bootstrap resampling. The green lines represent the residual spectrum, computed as the difference between the two stacks. *Top panel:* A zoom-in on the [NeV] emission lines complex from the full stacks shown in the bottom panel. While [NeV] λ 3346 in our stack is consistent with the noise level in the JWST stacked spectrum, [NeV] λ 3426 stands out at approximately six times that level. In this region, we applied a 0.08 offset to the M24 stack, as their flux is systematically below zero, likely due to their stacking method. *Bottom panel:* Full stacked spectra, highlighting the broader comparison between the two datasets. Major emission lines are labeled across the top for reference.

$\log N_{\text{H}}/\text{cm}^{-2} > 21.5$. Within this area, we highlight a more extreme subregion defined by $\log L_{\text{X}}/\text{erg s}^{-1} > 44.6$ and $\log N_{\text{H}}/\text{cm}^{-2} > 22$ with black dashed lines. This subregion is uniquely sampled by our dataset, whereas only three sources from other surveys fall within it. Two of these are from Ricci et al. (2017a): one (BAT ID 800) is only mildly obscured in the X-rays ($\log N_{\text{H}}/\text{cm}^{-2} \sim 22.3$) and it is not included in our sample because it is classified as a Seyfert 1; the other, NGC 6240 (BAT ID 841), consists of two nuclei (Komossa et al. 2003). In Ricci et al. (2017a), NGC 6240 was fitted as a single source due to the limited resolution of XMM-Newton data. However, when modeled separately, both nuclei have luminosities $\log L_{\text{X}}/\text{erg s}^{-1} < 44.6$ (Puccetti et al. 2016), placing them outside our defined luminosity selection threshold. The third source, from XMM-XXL (Liu et al. 2016), exhibits a broad $\text{H}\beta$ line in the optical spectrum (Rakshit et al. 2017), and it also has a relatively low SNR in the X-ray band (only ~ 60 photons detected in the 2–10 keV XMM-Newton PN spectrum). Liu et al. (2016) fitted the source with a double power-law model, reporting an unusually strong secondary component (scattering fraction $> 60\%$), which likely led to an overestimation of N_{H} (Tokayer et al. 2025). In any case, these are only a few sources located near the edge of the defined region, leaving this extreme parameter space uniquely well-sampled by our dataset.

7.2. On the $[\text{Ne v}]\lambda 3426$ emission line

We showed that our sample has an $85^{+5}_{-11}\%$ detection fraction for $[\text{Ne v}]\lambda 3426$, highlighting its utility as a tracer of AGN emission even in cases of heavy obscuration. This is particularly valuable for identifying AGNs where extreme absorption can completely suppress the X-ray emission. As such, $[\text{Ne v}]\lambda 3426$ becomes particularly relevant in light of recent high-redshift AGN candidates discovered with JWST, many of which exhibit weaker-than-expected or entirely undetected X-ray emission, even in stacked analyses (e.g., Ananna et al. 2024; M24). These include the so-called Little Red Dots (LRDs; e.g., Greene et al. 2024; Harikane et al. 2023; Kocevski et al. 2023, 2024; Matthee et al. 2024; Labbe et al. 2025), whose nature remains uncertain (e.g., Baggen et al. 2024; Kokubo & Harikane 2024). Proposed explanations for this X-ray weakness include heavy obscuration (e.g., Maiolino et al. 2024, 2025), super-Eddington accretion (e.g., Pacucci & Narayan 2024; Lambrides et al. 2024; Madau & Haardt 2024), or a weak or undeveloped X-ray corona (Juodžbalis et al. 2024). Still, the exact scenario remains unclear (e.g., Yue et al. 2024; M24; de Graaff et al. 2025). In this context, it is particularly puzzling that $[\text{Ne v}]\lambda 3426$ emission is often un-

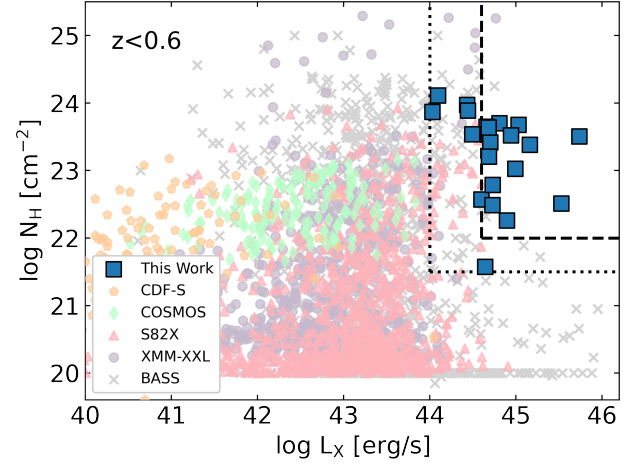


Figure 9. Intrinsic 2–10 keV luminosity (L_{X}) versus absorbing column density (N_{H}) for our sample (blue squares), compared with sources from other X-ray surveys with available spectral analyses. For consistency, only sources at $z < 0.6$ are shown to match the redshift range of our sample. Data from other surveys include: the Chandra Deep Field South (Liu et al. 2017; orange pentagons), COSMOS (Marchesi et al. 2016; green diamonds), Stripe 82X (Peca et al. 2023; pink triangles), XMM-XXL (Liu et al. 2016; purple circles), and BASS 70-month catalog (Ricci et al. 2017a; grey crosses). The region covered by our full sample is shown with black dotted lines, defined by $\log L_{\text{X}}/\text{erg s}^{-1} > 44$ and $\log N_{\text{H}}/\text{cm}^{-2} > 21.5$. Within this area, we highlight in black dashed lines a more extreme subregion, defined by $\log L_{\text{X}} > 44.6$ and $\log N_{\text{H}}/\text{cm}^{-2} > 22$, which is uniquely sampled by our sources but only poorly populated by previous surveys (see text for details).

detected in JWST-selected samples, including LRDs. R25 extend our findings by showing that roughly half of all BASS AGNs exhibit $[\text{Ne v}]\lambda 3426$ emission, and that this fraction can be as high as $\sim 70\%$ in high-SNR spectra. Importantly, they find that the detection fraction remains $\gtrsim 40\text{--}50\%$ across a wide range of luminosities, spanning from the high luminosities of our sample down to $\log L_{\text{X}}/\text{erg s}^{-1} \sim 42$, with no clear dependence on X-ray obscuration or Eddington ratio. This suggests that, even though the high-redshift, obscured AGNs identified by JWST seem to have lower bolometric luminosities and accretion rates (e.g., Yang et al. 2023; Lyu et al. 2024; M24; Maiolino et al. 2025) than those in our sample, we would still expect $[\text{Ne v}]\lambda 3426$ emission to be detectable in a larger fraction of those sources. However, only a few detections of $[\text{Ne v}]\lambda 3426$ have been reported so far (e.g., Chisholm et al. 2024; Labbe et al. 2024), in particular for narrow-line AGNs in deep fields such as CEERS (M24) and JADES (Eisenstein et al. 2023; Scholtz et al. 2025), challenging our current understanding of these objects.

Radiative transfer models by [McKaig et al. \(2024\)](#) further emphasize this puzzle. Their calculations indicate that [Ne v] λ 3426 is among the least dust-sensitive coronal lines, with its luminosity reduced by less than a factor of ten even under extreme conditions of super-solar metallicity and dusty gas, compared to a dust-free scenario. Thus, the general lack of detectable [Ne v] λ 3426 emission in JWST-selected AGNs suggests potentially different physical conditions compared to those typically found in local sources. An alternative explanation involves substantial, host-galaxy-scale obscuration by the interstellar medium (ISM). [Gilli et al. \(2022\)](#) showed that the observed increase in obscuration with redshift seen in X-ray surveys (e.g., [Aird et al. 2015](#); [Vijarnwanna et al. 2022](#); [Peca et al. 2023](#); [Pouliasis et al. 2024](#)) can be closely linked to the rising contribution of ISM-scale gas, which may reach column densities as high as $\log N_{\text{H}}/\text{cm}^{-2} \sim 24$ at $z \gtrsim 6$ (see also e.g., [Circosta et al. 2019](#); [D’Amato et al. 2020](#)). While this supports the obscured AGN scenario as a possible explanation for the observed X-ray weakness, such high column densities can also correspond to a significant amount of dust, which could attenuate emission lines like [Ne v] λ 3426.

Extinction measurements¹⁰ in our local sample and M24 are comparable, with median values of $A_V \approx 0.5$ – 0.6 (consistent also with [Harikane et al. 2023](#); [Maiolino et al. 2024](#)), suggesting similar total dust effects. At the same time, our local galaxies have higher stellar masses (median $\log M_*/M_\odot = 10.9$) and thus larger gas and dust reservoirs (e.g., [Santini et al. 2014](#); [Salvestrini et al. 2022](#)), compared to M24 and other JWST-selected hosts ($\log M_*/M_\odot < 10$; e.g., [Harikane et al. 2023](#); [Maiolino et al. 2024, 2025](#)). The similarity in extinction despite these mass differences suggests complex variations in dust and gas geometry, distributions, or dust grain properties, making direct comparisons challenging and beyond the scope of this work. However, since we matched our sample to M24 in [O III] λ 5007 (Section 6.2), we expect at least comparable ionization conditions between the samples. Indeed, although JWST hosts have typically lower stellar masses and metallicities (e.g., [Maiolino et al. 2024](#); [Curti et al. 2024](#)), these effects are mitigated by the [O III] λ 5007 matching, as stellar mass correlates with metallicity (e.g., [Tremonti et al. 2004](#); [Curti et al. 2020](#)) and higher metallicity weakens [O III] λ 5007 emission efficiency (e.g., [Kewley & Dopita 2002](#); [Maiolino et al. 2008](#)). Therefore, the local and JWST stacks compared

in this work are expected to have similar ionization conditions, making the observed differences in [Ne v] λ 3426 detectability surprising.

Future and upcoming efforts using JWST data (e.g., with NIRSpec) will play a crucial role in resolving these open questions, providing deeper insights into the intrinsic properties and nature of these intriguing high-redshift AGNs. For example, [Cleri et al. \(2023b\)](#) demonstrated that [Ne v] λ 3426 can serve as a valuable diagnostic tool for identifying AGN activity in high-redshift line-emitting galaxies, particularly where traditional line-ratio diagnostics may be less effective. Indeed, recent JWST-based studies (e.g., [Chisholm et al. 2024](#); [Scholtz et al. 2025](#)) have successfully relied on [Ne v] λ 3426, along with other high-ionization emission lines, to argue for an AGN-dominated nature in several galaxies at $z > 5$. In this regard, an upcoming publication by [Trakhtenbrot et al. \(submitted\)](#) further explores the utility of [Ne v] λ 3426 to infer AGN luminosities, and thus constraints on early black hole growth, by leveraging available JWST data for high-redshift galaxies.

7.3. Future X-ray studies

This work underscores the importance of local AGNs for understanding the nature of X-ray emission and its connection to multi-wavelength properties. It also highlights the need for future facilities, such as the Advanced X-ray Imaging Satellite (*AXIS*; [Reynolds et al. 2024](#)) and *NewAthena* ([Nandra et al. 2013](#); [Barret et al. 2020](#)), to replicate such detailed analyses at higher redshifts. In what follows, we discuss the capabilities of the two instruments based on the most up-to-date planned exposures, survey coverages, and data products available from both teams at the time of writing¹¹. For *NewAthena*, we specifically focus on the Wide Field Imager (WFI; [Meidinger et al. 2017](#)), which is expected to play a central role in future X-ray surveys.

First, we quantify how many sources similar to those in our sample will be detected. To do so, we used the mock catalog of [Marchesi et al. \(2020\)](#), which is based on the luminosity functions of [Gilli et al. \(2007\)](#) for $z < 3$ and [Vito et al. \(2014, 2018\)](#) at higher redshifts. Assuming the *AXIS* limiting fluxes ([Reynolds et al. 2024](#)) of $2.5 \times 10^{-17} \text{ erg s}^{-1} \text{ cm}^{-2}$ for the Intermediate survey (7000 arcmin^2) and $4 \times 10^{-18} \text{ erg s}^{-1} \text{ cm}^{-2}$ for the Deep survey (450 arcmin^2), we predict the detection of a total of ~ 150 luminous ($\log L_X/\text{erg s}^{-1} > 44.6$) and obscured ($\log N_{\text{H}}/\text{cm}^{-2} > 22$) AGNs. Of these, about 100 sources will be at $z < 3$, and approximately 50 at $z > 3$.

¹⁰ Computed as in M24, using the Balmer decrement, assuming Case B recombination, and adopting the [Calzetti et al. \(2000\)](#) extinction law with $R_V = 3.1$.

¹¹ <https://blog.umd.edu/axis/>;
<https://www.mpe.mpg.de/ATHENA-WFI/>

For *NewAthena*, we used the three planned surveys described in [Cruise et al. \(2025\)](#), with limiting fluxes of $\sim 5 \times 10^{-17}$ erg s $^{-1}$ cm $^{-2}$ for the deeper ~ 31 deg 2 survey, and 4×10^{-17} and 7×10^{-17} erg s $^{-1}$ cm $^{-2}$ for the wider surveys of approximately 180 and 370 deg 2 , respectively. Using these sensitivities and areas, we predict the detection of ~ 22700 luminous and obscured AGNs, with about 14800 sources at $z < 3$ and 7900 at $z > 3$. This significant increase reflects the spatial rarity of these objects (e.g., [Ueda et al. 2014](#); [Aird et al. 2015](#); [Peca et al. 2023](#)), which can only be captured in large numbers through wide-area surveys.

While these predictions highlight the detection power of future surveys, detailed spectral analysis requires higher photon statistics than simple source detection. To evaluate the ability of *AXIS* and *NewAthena* to recover key spectral parameters for sources like those in our sample, we performed dedicated simulations based on one of our best-studied AGNs, 2MASX J01290761-6038423 (BAT ID 80), which has the largest number of spectra in our dataset. Its best-fit spectral model at $z = 0.203$ was redshifted to $z = 1, 3$, and 5 , while preserving the intrinsic luminosity (e.g., [Peca et al. 2024a](#); [Tokayer et al. 2025](#)). We considered both typical single-pointings exposures of 20 ks and 125 ks, as well as planned survey exposures of 335 ks and 7 Ms for *AXIS*, and 200 ks for *NewAthena*. On-axis responses and backgrounds were used for single-pointings, while field averaged files were used for survey cases. Simulations were limited to $z \leq 5$, as with these exposures *NewAthena* does not meet the minimum count threshold (see below) at higher redshifts, and *AXIS* is constrained by the low number of high-luminosity sources expected in its planned surveys. At $z = 3$ and 5 , we excluded the thermal *apec* component, as it shifts out of the observable band. Simulated spectra were analyzed over the 0.5–10 keV band using the same methods discussed in Section 4, with 100 realizations per redshift–exposure combination. For spectra with less than 200 counts (down to 70 total counts, following [Peca et al. 2023](#)), we grouped at a minimum of 5 counts per bin and applied Cash statistics ([Cash 1979](#)).

Results are shown in Figure 10. At the same redshift of the simulated source ($z = 0.203$), both observatories perform well across all exposures. At higher redshifts, *NewAthena* provides better results for single-pointing exposures, while *AXIS* is limited to $z = 1$ under the same conditions. For the deeper exposures in its planned surveys, *AXIS* achieves excellent parameter recovery out to $z = 5$, particularly in the 7 Ms case, benefiting from deep exposure, low background, and minimal off-axis degradation. In contrast, *NewAthena*’s 200 ks survey

offers only marginal improvement over its 125 ks single-pointing, likely due to background and off-axis effects, but still delivering measurements up to $z = 5$.

Overall, this comparison shows the complementary strengths of the two missions: *NewAthena* is well suited for detailed studies of bright sources with short to moderate exposures, while *AXIS* is optimized for deep surveys of faint AGNs due to its low background and stable off-axis performance. Together, these complementary capabilities will expand the redshift and luminosity range over which detailed spectral studies can be conducted. In this context, the best-fit models derived in this work offer valuable templates for testing and interpreting future observations of high-luminosity, obscured AGNs with both current and next-generation X-ray missions.

8. SUMMARY

In this work, we analyzed a sample of 21 highly luminous X-ray AGNs ($\log L_X/\text{erg s}^{-1} > 44.6$), optically classified as Seyfert 1.9 and 2. We conducted a detailed spectral analysis using *NuSTAR*, *XMM-Newton*, *Chandra*, and *Suzaku* data, incorporating multi-epoch observations where available. This sample probes a region of the luminosity–absorption column density parameter space that is covered almost exclusively by the BASS survey. This high-quality dataset enabled physically motivated modeling to constrain both physical and geometrical parameters, with extensive MCMC analyses used to fully explore the parameter space. In addition, we analyzed the available optical spectra, with particular focus on the detection and interpretation of the [Ne v] $\lambda 3426$ emission line.

Our results can be summarized as follows:

- We found a weak correlation ($2 < \sigma < 3$) between the photon index (Γ) and Eddington ratio (λ_{Edd}), consistent with prior studies. Our results confirm that the relation extends to highly luminous and obscured AGNs;
- Half (6/12) of the sources with black hole mass estimates lie in the absorbing column density (N_{H})- λ_{Edd} “forbidden” region. This suggests that highly luminous, obscured AGNs are likely to be caught in a transitional phase, where radiation from the central engine is beginning to clear the surrounding obscuring material. In support of this scenario, we found that sources within and near this region exhibit a higher incidence of N_{H} variability and outflows compared to those located outside this regime;

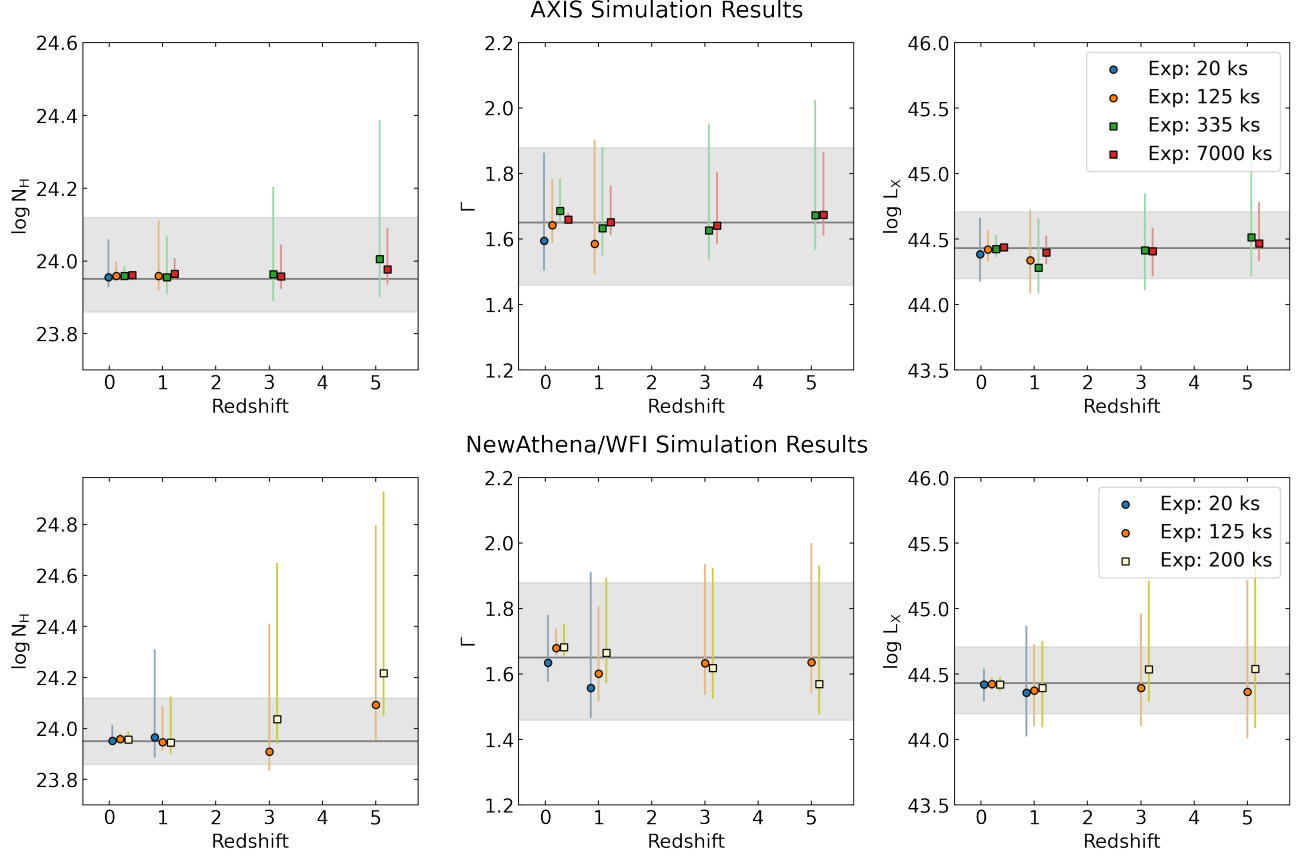


Figure 10. Comparison of column density ($\log N_H$; left), photon-index (Γ ; center), and intrinsic luminosity ($\log L_X$; right) as a function of redshift for simulated observations with *AXIS* (top panels) and *NewAthena* (bottom panels). Simulations are based on the best-fit model of 2MASX J01290761-6038423 (BAT ID 80, $z = 0.203$), redshifted to $z = 1, 3$, and 5 , and modeled for a range of exposure times. Single-pointing-like exposures of 20 ks (blue) and 125 ks (orange) were simulated for both missions. Survey-depth exposures include 335 ks (green) and 7 Ms (red) for *AXIS*, and 200 ks (yellow) for *NewAthena*, consistent with its planned deep survey. Circles represent single-exposure simulations, while squares indicate survey-depth cases. Points show the median values from the MCMC posteriors; error bars represent the 5th–95th percentile range. Points are slightly offset in redshift for visual clarity. Horizontal grey lines mark the original best-fit values from the real source, with shaded regions indicating their associated uncertainties. All simulations used field-of-view averaged response files and realistic background levels for each mission. Points are omitted where simulations failed to produce sufficient counts for spectral analysis.

- The covering factors in our sample are consistent with those of Compton-thin AGNs and are systematically higher than values typically found for Compton-thick sources. This supports a scenario in which Compton-thin material extends to larger covering factors than Compton-thick obscuration;
- The measured iron $K\alpha$ equivalent widths (EWs) reflect the observational trend of increasing EW with higher column density, as the continuum becomes more suppressed relative to the line emission. We also observe a flattening of this trend at $\log N_H/\text{cm}^{-2} < 23$, consistent with the iron line originating from material in the inner edge of the torus or from the broad line region. Additionally, our results are consistent with the Bald-

win effect, with higher-luminosity sources showing lower EWs;

- For sources with available multi-epoch spectroscopy, we performed a variability analysis for our sample. We found that 11 out of 13 sources ($85^{+5}_{-15}\%$) exhibit variability either in flux ($10/13$, $77^{+8}_{-15}\%$) or N_H ($4/12$, $33^{+15}_{-10}\%$). These results are in excellent agreement with recent findings on lower-luminosity ($\log L_X/\text{erg s}^{-1} \sim 43$), Compton-thin AGNs.
- For 20/21 sources with adequate optical spectroscopy, we measured the $[\text{Ne V}]\lambda 3426$ emission line and found a detection rate of $85^{+5}_{-11}\%$ (17/20). This confirms that the $[\text{Ne V}]\lambda 3426$ emission line

is a strong tracer of AGN activity even in cases of heavy obscuration, in agreement with previous studies;

- We stacked the optical spectra in our sample and compared the result with stacked spectra of JWST-selected, narrow-line AGNs at $2 \lesssim z \lesssim 9$, which show X-ray weakness. We found that while $[\text{Ne V}]\lambda 3426$ is strongly detected in our local AGN stack, it is absent in the high-redshift JWST stack. Given that we matched samples in $[\text{O III}]\lambda 5007$ to ensure comparable ionization conditions, this discrepancy suggests that the nature of X-ray weak JWST AGNs at high redshift may be more complex than that of their local analogs, potentially involving different physical conditions, evolutionary stages, or environments in which $[\text{Ne V}]\lambda 3426$ is more difficult to detect.
- We performed simulations using *AXIS* and *NewAthena* to simulate how sources similar to those in our sample would appear with next-generation X-ray telescopes. Using our best-fit models as templates, we show that both missions will offer complementary capabilities and will be essential for extending this study to higher redshifts.

Our results highlight two key considerations. First, they illustrate how high-quality X-ray observations are necessary for performing detailed spectral modeling, which can be challenging at higher redshifts with current instrumentation. Thus, the models and results presented here can serve as valuable benchmarks for future studies with planned next-generation X-ray facilities, such as *AXIS* and *NewAthena*. Second, local AGN studies that combine detailed X-ray spectroscopy with multi-wavelength data are crucial for interpreting recent JWST discoveries of high-redshift sources whose nature remains ambiguous. In particular, our analysis of the $[\text{Ne V}]\lambda 3426$ emission line demonstrates how local AGNs can serve as useful laboratories, providing valuable insights into the physical conditions and possible obscuration properties of AGNs in the early Universe.

Facilities: *NuSTAR*, *XMM-Newton*, *Chandra*, *Suzaku*

Software: NuSTARDAS v2.1.4¹², SAS v21.0.0 (Gabriel et al. 2004), CIAO v4.16.0 (Fruscione et al. 2006), PyXspec v2.1.3 (Gordon & Arnaud 2021; Arnaud

1996), TOPCAT v4.94.9 (Taylor 2005), Astropy v7.0.1 (Astropy Collaboration et al. 2013, 2018, 2022); Pandas v2.1.4 (Wes McKinney 2010; pandas development team 2020), SciPy v1.15.2 (Virtanen et al. 2020), Matplotlib v3.10.1 (Hunter 2007), linmix v0.1.0 (Kelly 2007).

ACKNOWLEDGMENTS

We acknowledge the anonymous referee for the valuable comments that improved the quality of the paper. A.P. acknowledges D. Costanzo for all the support over the years. A.P. and M.K. acknowledge support from NASA through ADAP award 80NSSC22K1126 and NuSTAR grants 80NSSC22K1933 and 80NSSC22K1934. K.O. acknowledges support from the Korea Astronomy and Space Science Institute under the R&D program (Project No. 2025-1-831-01), supervised by the Korea AeroSpace Administration, and the National Research Foundation of Korea (NRF) grant funded by the Korea government (MSIT) (RS-2025-00553982). B.T. acknowledges support from the European Research Council (ERC) under the European Union’s Horizon 2020 research and innovation program (grant agreement number 950533) and from the Israel Science Foundation (grant number 1849/19). C.R. acknowledges support from Fondecyt Regular grant 1230345, ANID BASAL project FB210003 and the China-Chile joint research fund. A.T. acknowledges financial support from the Bando Ricerca Fondamentale INAF 2022 Large Grant “Toward an holistic view of the Titans: multi-band observations of $z > 6$ QSOs powered by greedy supermassive black holes”. I.M.C. acknowledges support from ANID programme FONDECYT Postdoctorado 3230653. M.S. acknowledges financial support from the Italian Ministry for University and Research, through the grant PNRR-M4C2-I1.1-PRIN 2022-PE9-SEAWIND: Super-Eddington Accretion: Wind, Inflow and Disk - F53D23001250006 - NextGenerationEU.

¹² <https://heasarc.gsfc.nasa.gov/docs/nustar/analysis/>

APPENDIX

A. NOTES ON INDIVIDUAL SOURCES

The baseline model described in Section 4 was adjusted for single sources, where necessary, according to the model selection procedure described in Section 5.1. A second *apec* model was included to improve the fit at soft energies for LEDA2816387 (BAT ID 199). For 2MASS J02162672+5125251 (BAT ID 119), LEDA 511869 (BAT ID 714), PKS 1549-79 (BAT ID 787), GALEXMSC J025952.92+245410.8 (BAT ID 1248), and GALEXASC J063634.15+591319.6 (BAT ID 1296), we turned off the *apec* emission, as the secondary power-law component adequately accounted for the soft photons from either *Chandra* and XMM-*Newton* spectra, without the need for an additional component. For 2MASS J17422050-5146223 (BAT ID 1515), we turned off both the *apec* and scattering components, as the primary power-law emission alone provides an adequate fit to the soft X-ray spectrum. Indeed, the source is unobscured in the X-rays, with an upper limit on N_{H} of $1.2 \times 10^{21} \text{ cm}^{-2}$, consistent with the high Galactic column density in that direction.

For RBS 2043 (BAT ID 1204) and 2MASX J09003684+2053402 (BAT ID 1346), which reside at the center of the Phoenix and Z2089 galaxy clusters, respectively, we used two *apec* components to model the cluster thermal emission. Specifically, we used sub-solar abundances (~ 0.33 and ~ 0.76 for the two components) as described by Ueda et al. (2013) in a detailed analysis of the Phoenix cluster. The temperatures of the components, shown in Table 2, agree with their results. We used the same approach for Z2089, as we found a satisfactory fit using the same abundances. Here, the temperatures of the thermal components are in agreement with the thermal profile reported in Giles et al. (2017).

For RBS 2043 (BAT ID 1204), we found a larger scattering fraction ($\sim 14\%$) than the other sources in our sample. This value remains consistent, within the uncertainties, with the results of Ricci et al. (2017a). While such a high scattering fraction is in line with findings for high luminosity AGNs in X-ray surveys (e.g., Ricci et al. 2017a; Peca et al. 2023), we acknowledge that it may be affected by the cluster thermal emission, as both the components cover the same energy range.

In the XMM-*Newton* 2021 observation of 2MASX J01290761-6038423 (BAT ID 80), we used only the PN data, as the MOS cameras lacked good coverage in the detector chips where the source was observed after background filtering. We utilized only the PN camera also

for 2MASS J02162672+5125251 (BAT ID 119) and PKS 1549-79 (BAT ID 787), as both MOS cameras were affected by pile-up issues.

The *Chandra* observation used for BAT ID 1248 is contained in the Chandra Data Collection DOI: [10.25574/cdc.408](https://doi.org/10.25574/cdc.408)

The best-fit unfolded spectra, models, and residuals are presented in Figures 16, 17, and 18. For sources with multi-epoch observations, the temporal evolution of C_{AGN} , N_{H} , and the observed 2–10 keV flux (see Section 5.6) is shown in Figures 19, 20, and 21. To maintain graphical clarity and flow in the text, all these figures are placed at the end of the paper.

We also show in this Section that the N_{H} values derived from different torus models presented in Section 4 are consistent within the uncertainties. This is shown in Figure 11, which compares the fitted column densities for individual sources across models. All values agree within 0.2 dex when considering the uncertainties, with no significant systematic offsets found. This indicates that the inferred obscuration levels are consistent and not significantly dependent on, or biased by, the choice of torus model.

B. COMPARISON WITH BASS DR1 AND DR2

In this appendix, we compare the results from our spectral analysis with those of the BASS DR1 and DR2, which are based on the *Swift*/BAT 70-month catalog. Specifically, we compare N_{H} , L_{X} , and Γ with Ricci et al. (2017a), and λ_{Edd} with Koss et al. (2022b), as shown in Figure 12 and 13, respectively. The primary difference between this work and Ricci et al. (2017a) is the inclusion of *NuSTAR* observations, which were unavailable for the earlier analysis. Additionally, there are differences in spectral modeling: while Ricci et al. (2017a) employed a compilation of simple to complex models, they did not use the self-consistent AGN torus models adopted here.

Figure 12 shows the comparison for the 11 sources analyzed both in this work and Ricci et al. (2017a). The remaining sources are part of the BASS 105-month catalog only, and one source (BAT ID 494) has not been fitted in Ricci et al. (2017a). Overall, our results are consistent with those of Ricci et al. (2017a), yet with some scatter. Specifically, for L_{X} , all sources are within 0.5 dex, while for N_{H} all but two sources are within 0.5 dex, with 10/11 within one dex. A notable outlier in the N_{H} measurements is PKS 1549-79 (BAT ID 787), which Ricci et al. (2017a) identified as unobscured. However,

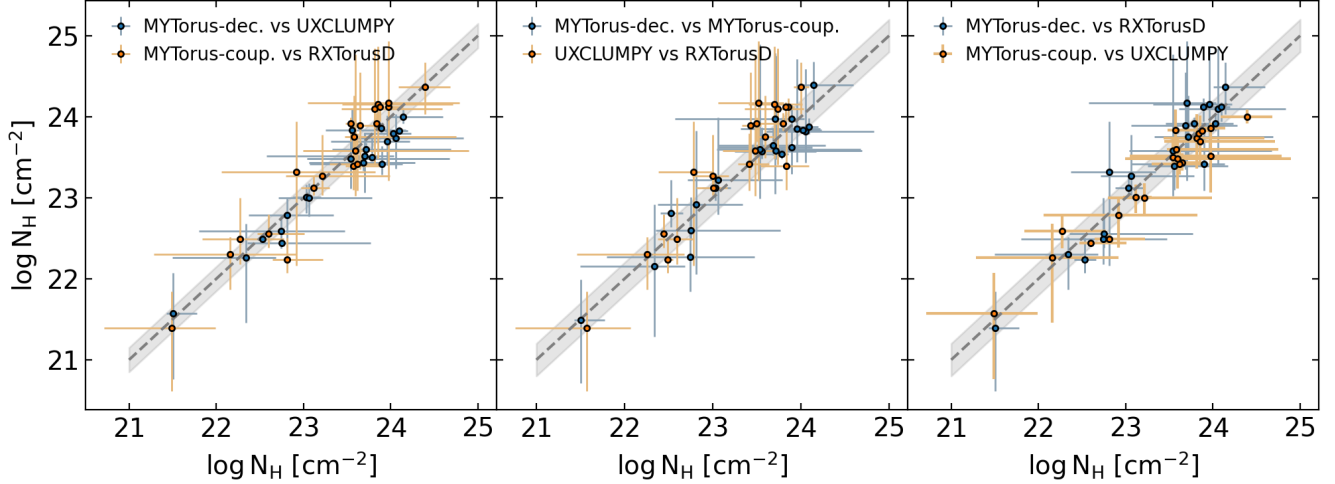


Figure 11. Comparison of N_H values derived using different torus models for the sources in our sample. Each panel shows pairwise comparisons between models, with the 1:1 relation indicated by a dashed grey line and a shaded region representing ± 0.2 dex deviation. *Top panel:* MYTorus-decoupled vs. UXCLUMPY (blue) and MYTorus-coupled vs. RXTorusD (orange). *Middle panel:* MYTorus-decoupled vs. MYTorus-coupled (blue) and UXCLUMPY vs. RXTorusD (orange). *Bottom panel:* MYTorus-decoupled vs. RXTorusD (blue) and MYTorus-coupled vs. UXCLUMPY (orange). Across all comparisons, the N_H values are consistent within the uncertainties, with no significant offsets.

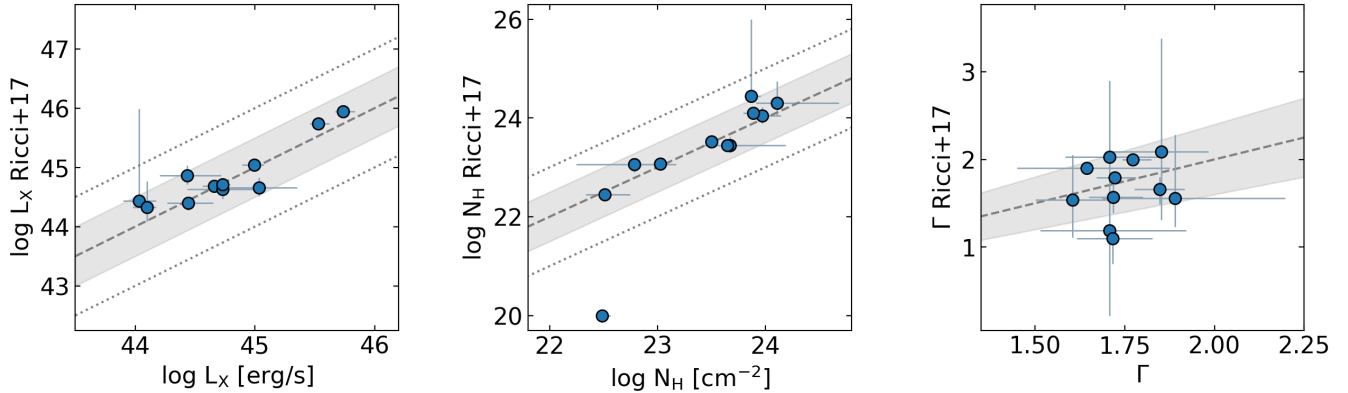


Figure 12. Comparison of L_X , N_H , and Γ (left, middle, and right panels, respectively) with Ricci et al. (2017a) for the 11 sources that are in both samples. The grey dashed lines represent the 1:1 relations. The shaded grey areas in the left and middle panels indicate the 0.5 dex deviation, and the grey dotted lines show the 1 dex deviation. For the right panel, the grey shaded area shows the 20% deviation from the 1:1 relation. There is an overall agreement for these parameters (details in text), with differences primarily due to the inclusion of *NuSTAR* data, different modeling, and variability.

our analysis classifies it as obscured. Our N_H measurement for this source is consistent with Tombesi et al. (2014), who analyzed the *XMM-Newton* data. We note that Ricci et al. (2017a) modeled this source based on its optical classification as a blazar, whereas more recent SED analyses incorporating optical and radio data now classify it as a Seyfert galaxy (Koss et al. 2022b and references therein). This, together with hints of variability (see Section 5.6) and the inclusion of *NuSTAR* data, might explain the difference in the results. For Γ , 9 out of 11 sources lie within a 20% scatter, and one of the two remaining sources is consistent with this range

when uncertainties are considered. The only clear outlier is ESP 39607 (BAT ID 32), which shows a very flat photon index of $\Gamma \sim 1.1$ in Ricci et al. (2017a). We note that the results from Peca et al. (2025), which focus specifically on this source, align well with the value obtained in this work.

To place our sample in the context of previous BASS population studies, we compare with the results of Ricci et al. (2017a) by selecting sources similar to those in our sample. Specifically, selecting Seyfert 1.9 and 2 AGNs with intrinsic (absorption-corrected) 2–10 keV luminosities above 10^{44} erg s $^{-1}$, the observed obscured fraction

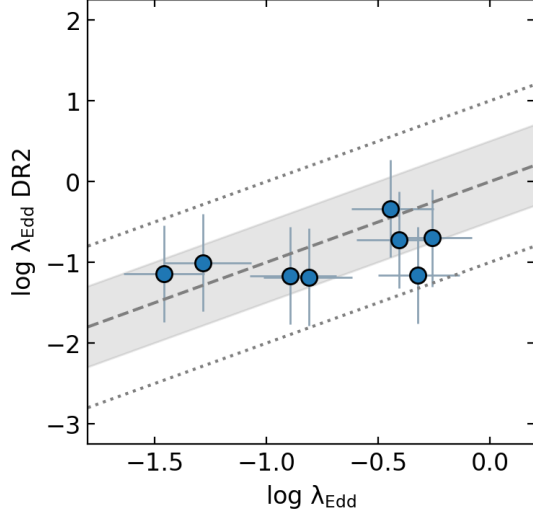


Figure 13. Comparison of λ_{Edd} with an updated version of the BASS DR2 catalog Koss et al. (2022b) (see text). λ_{Edd} has been estimated for 10 sources in this work. The grey dashed line represents the 1:1 relation, the shaded grey area indicates a 0.5 dex deviation, and the grey dotted lines show a 1 dex deviation. There is an overall agreement for these parameters (details in text), with differences primarily due to the inclusion of *NuSTAR* data, different modeling, and variability.

(i.e., sources with $N_{\text{H}} > 10^{22} \text{ cm}^{-2}$) in Ricci et al. (2017a) is $85^{+4}_{-7}\%$. In our sample, the corresponding obscured fraction is $95^{+2}_{-9}\%$, i.e., in agreement within the uncertainties.

Figure 13 shows a comparison of λ_{Edd} for eight sources for which this quantity was available both in our analysis and in BASS DR2. All the sources are within one dex, with 7/8 within 0.5 dex. The remaining source is consistent with this range when accounting for uncertainties. Differences in λ_{Edd} primary stem from updated black hole masses and new X-ray luminosities derived in this work. Overall, our results are in good agreement with prior estimates in BASS data releases.

C. OPTICAL/INFRARED AGN DIAGNOSTICS

This section presents the infrared and optical classifications for our sample. Figure 14 shows AGN selection criteria from Stern et al. (2012) and Mateos et al. (2012) based on *WISE* (Wright et al. 2010; Mainzer et al. 2011) mid-infrared color diagnostics. We matched our sample to the AllWISE catalog¹³ using a $1''$ matching radius, finding counterparts for all the sources. Vega magnitudes are used. All but one source satisfy the AGN selection criteria from either Stern et al. (2012) or Ma-

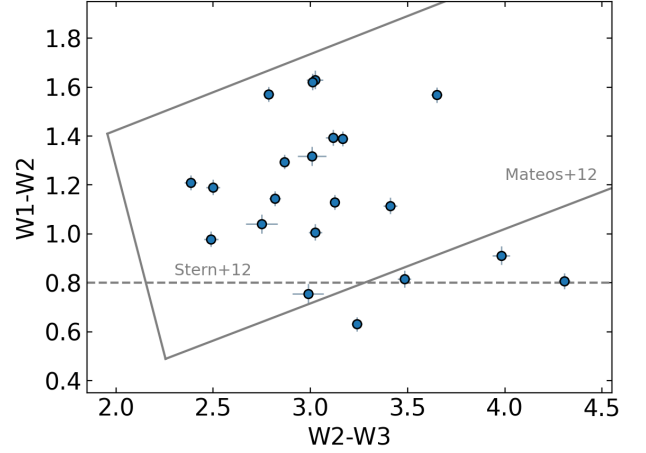


Figure 14. Color-color classification using *WISE* *W1*, *W2*, and *W3* color cuts. Sources within the region outlined by the solid grey lines (Mateos et al. 2012) or above the dashed grey line (Stern et al. 2012) are classified as AGN. One source falls outside both criteria but lies very close to the boundaries; see also Ichikawa et al. (2017).

teos et al. (2012). The single outlier (BAT ID 714) lies just outside both boundaries; however, Ichikawa et al. (2017) showed that *WISE*-based classifications can depend on luminosity, suggesting this source can still be safely classified as AGN.

Figure 15 shows the standard optical emission-line diagnostic diagrams (BPT; Baldwin et al. 1981; Veilleux & Osterbrock 1987). Emission line measurements for BASS DR2 sources are taken from Oh et al. (2022), while those for the remaining sources are derived from our new optical spectra (see Section 2.3) using the same methodology. The line diagnostics could be performed for 20 out of 21 sources, with the remaining source (BAT ID 119) lacking sufficient SNR for reliable line measurements. Specifically, we obtained measurements for 19 sources in the $[\text{N II}]\lambda 6583/\text{H}\alpha$ and $[\text{S II}]\lambda\lambda 6716, 6731/\text{H}\alpha$ versus $[\text{O III}]\lambda 5007/\text{H}\beta$ diagrams, and for 18 sources in the $[\text{O I}]\lambda 6300/\text{H}\alpha$ versus $[\text{O III}]\lambda 5007/\text{H}\beta$ diagram. As shown in the figure, our sources consistently fall within the Seyfert region across all three diagnostics. All sources described in this work are, therefore, classified as AGNs based on their infrared properties and exhibit Seyfert-like characteristics in the optical diagnostics.

¹³ <https://wise2.ipac.caltech.edu/docs/release/allwise/>

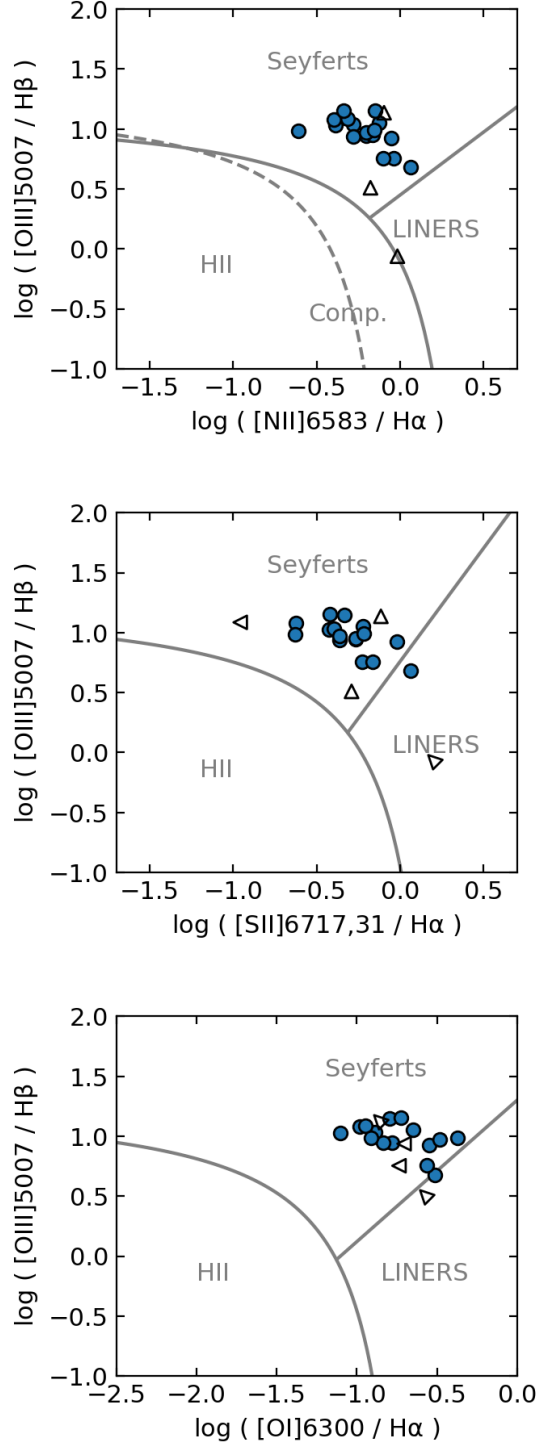


Figure 15. Optical classification from emission line diagnostic diagrams (Baldwin et al. 1981; Kewley et al. 2001; Kauffmann et al. 2003; Kewley et al. 2006; Schawinski et al. 2007). Upper limits are represented by empty markers oriented toward the limit. The uncertainties on emission line measurements are dominated by a systematic error of <0.1 dex (Koss et al. 2022b). All the sources without upper limits are classified as Seyfert in at least one diagnostic diagram.

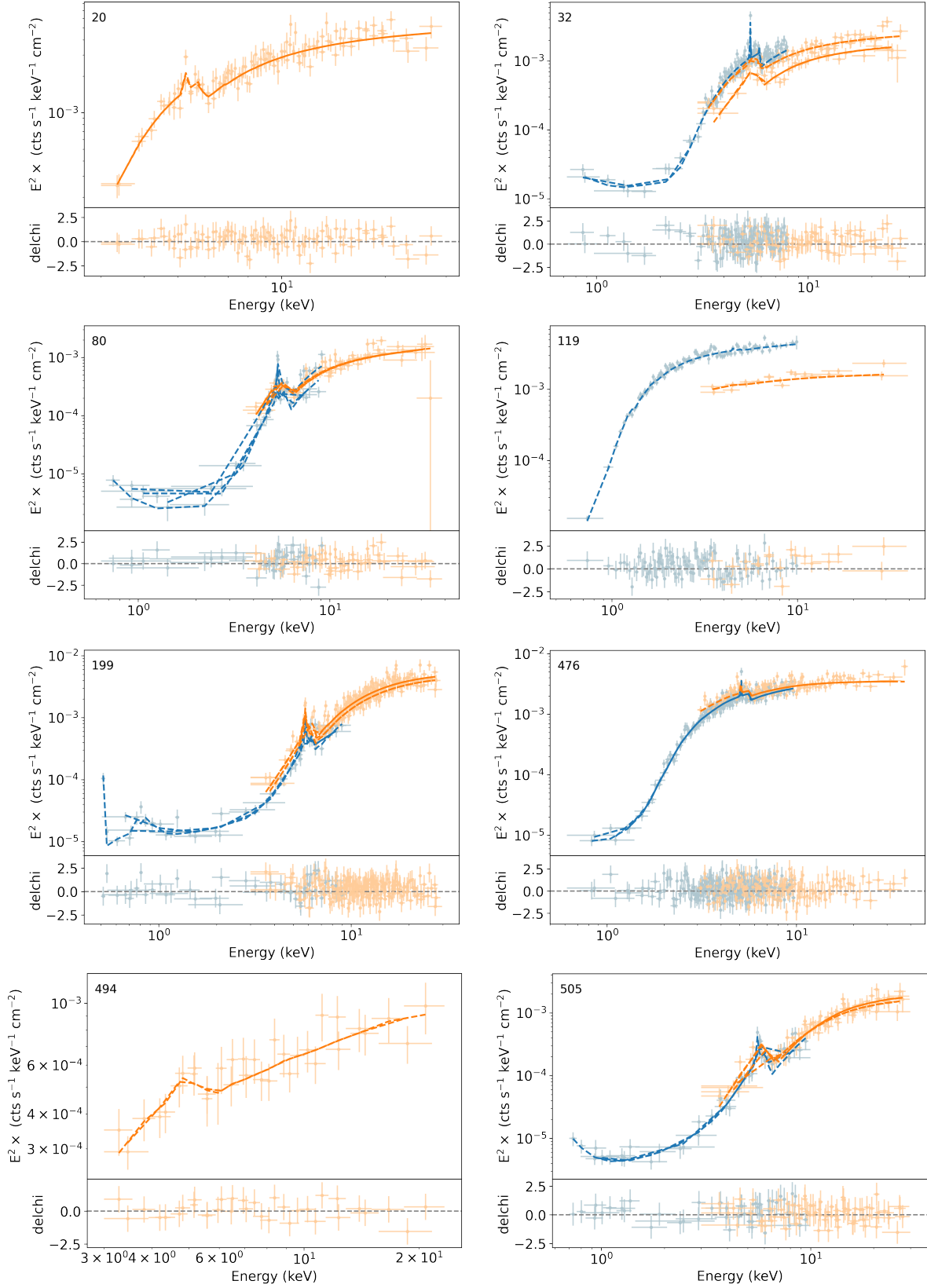


Figure 16. Best-fit unfolded spectra of the sources analyzed in this work. The corresponding models are shown as dashed lines, and the BAT ID is indicated in the top-left corner of each panel. For clarity, *NuSTAR* spectra are shown in orange, while soft X-ray spectra (*XMM-Newton*, *Chandra*, and *Suzaku*) are shown in blue. Residuals are displayed in the lower subpanels. All spectra have been rebinned for visual clarity.

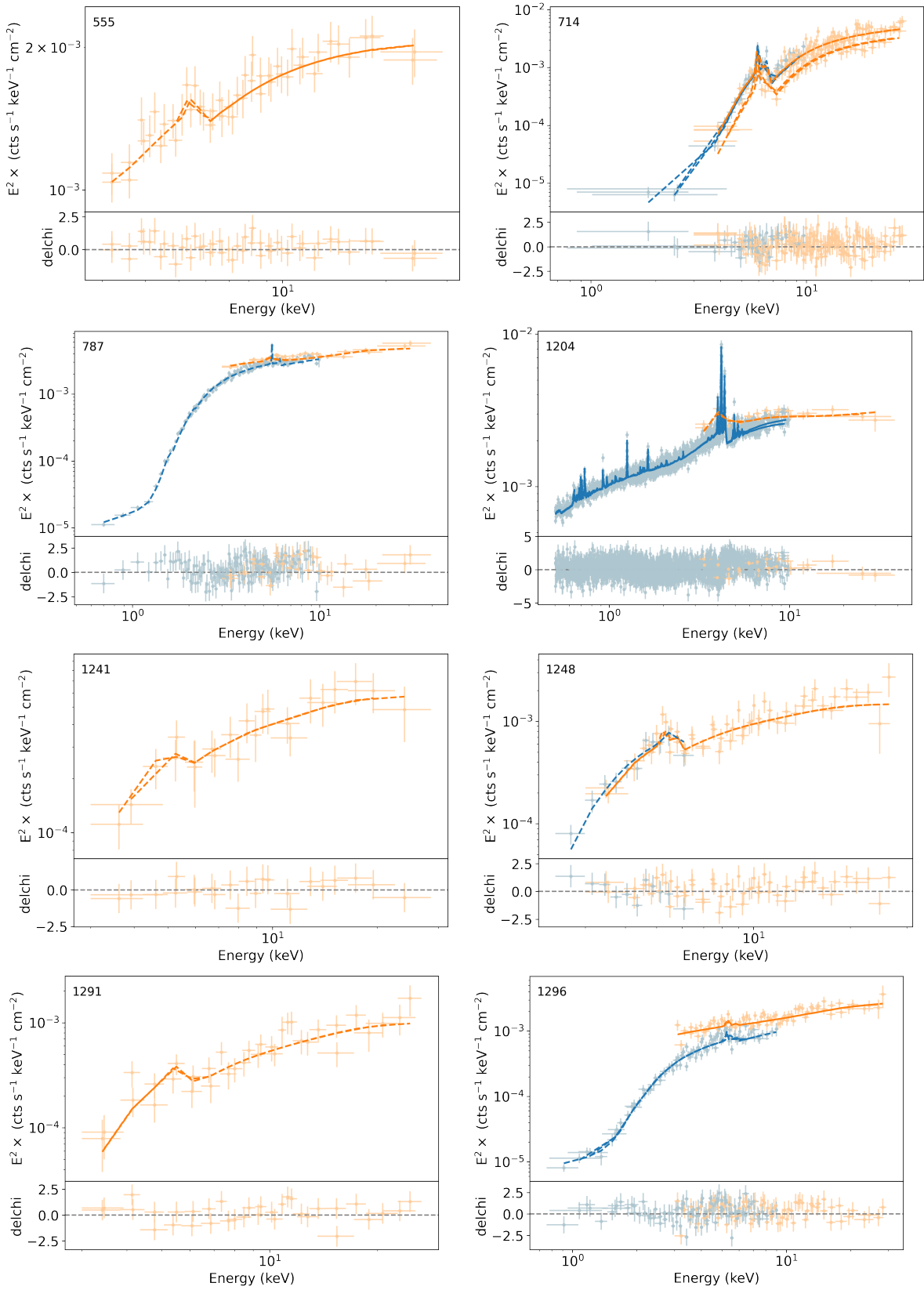
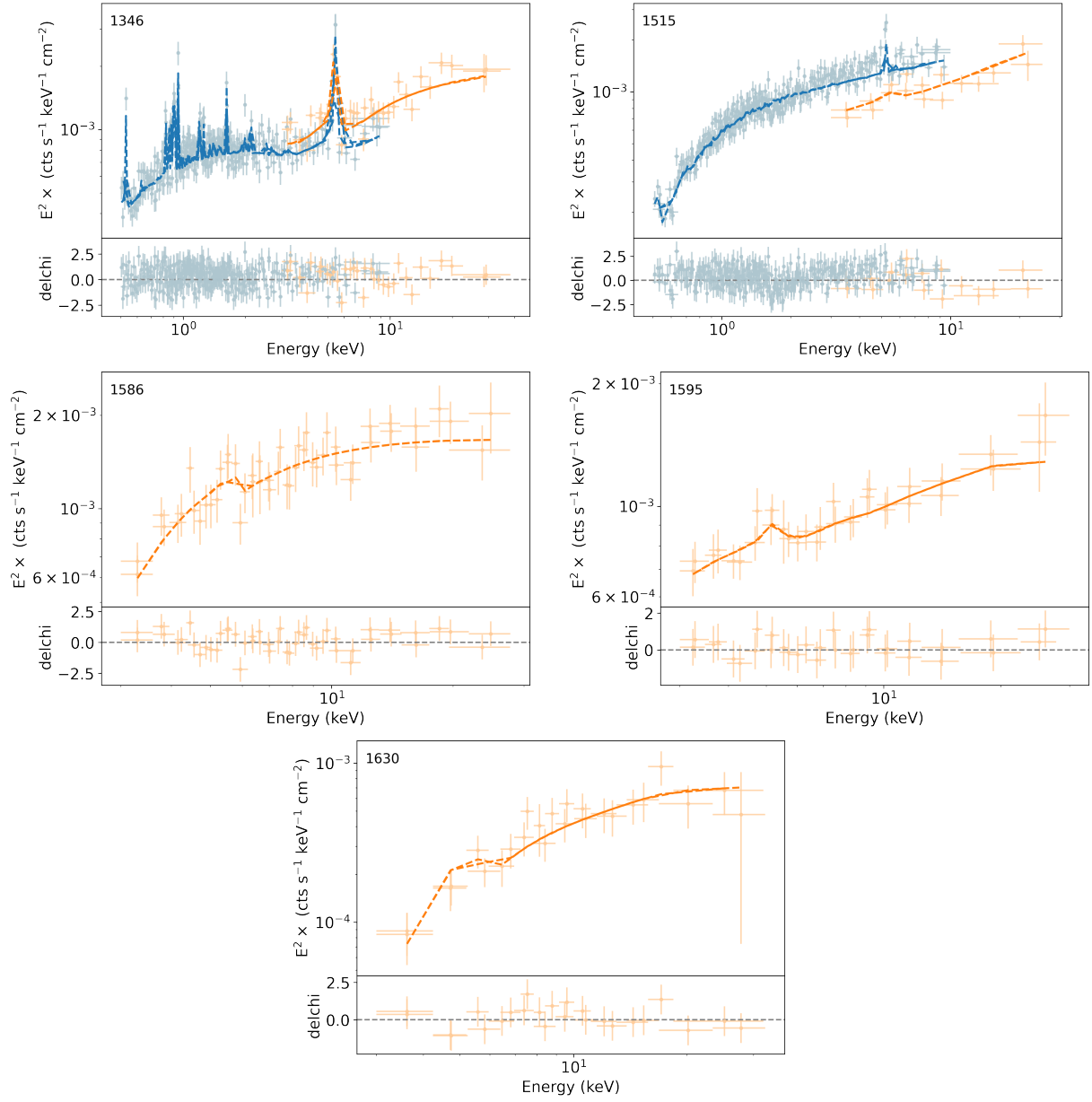


Figure 17. Same as Figure 16.

**Figure 18.** Same as Figure 16.

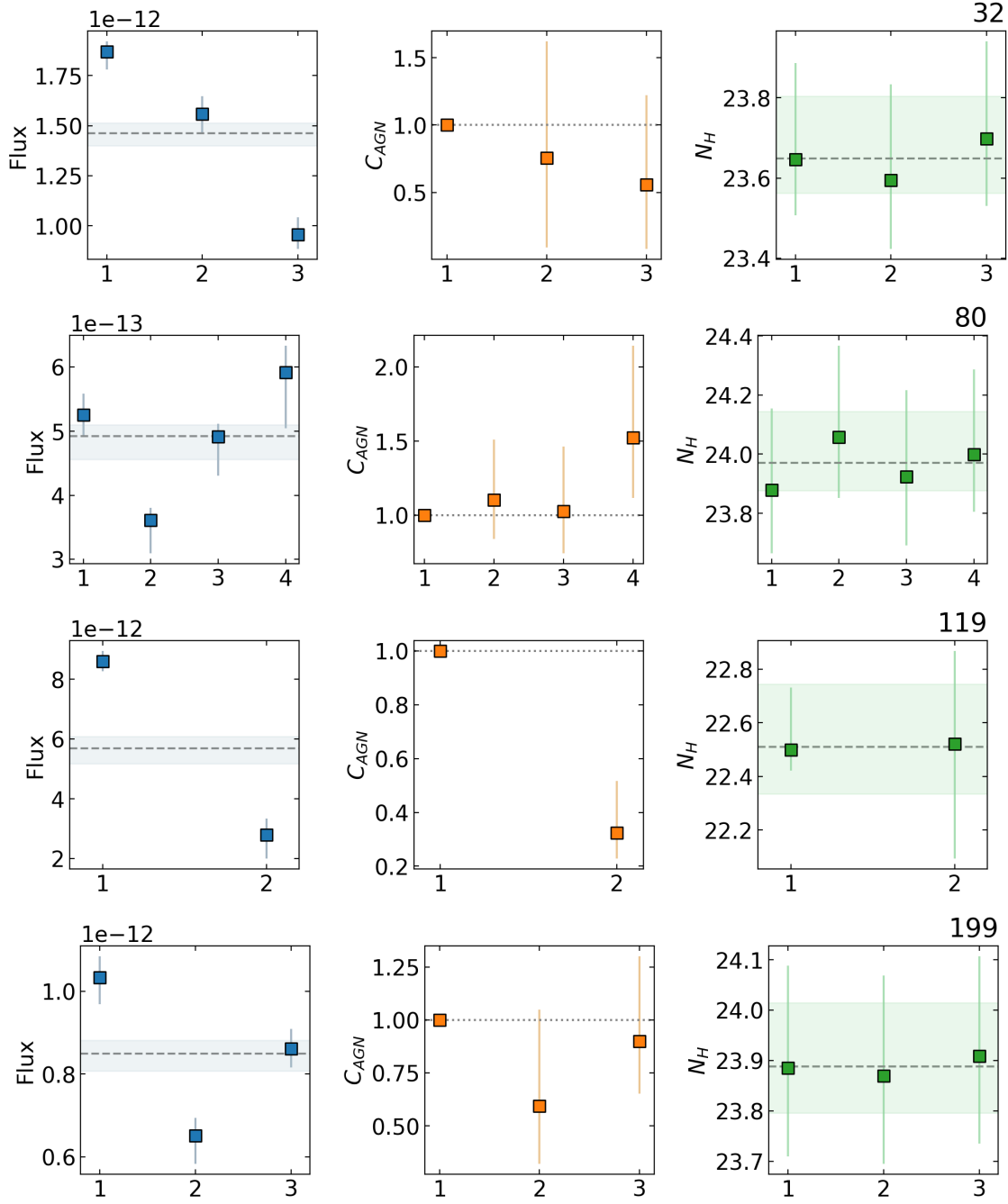


Figure 19. Observed 2–10 keV best-fit flux in erg/s/cm^2 (left panels), the constant C_{AGN} accounting for intrinsic variability (middle panels), and logarithm of line-of-sight N_H in cm^{-2} (right panels) for sources with multi-epoch observations. BAT IDs are indicated in the top-right of each panel. Horizontal lines in the flux and N_H panels indicate the average values, with shaded regions representing the propagated uncertainties. For C_{AGN} , the horizontal line and the first data point correspond to $C_{AGN} = 1$, fixed for the first epoch in the simultaneous fitting procedure. Observations are shown in temporal order (as listed in Table 4), though the x-axis spacing is not to scale.

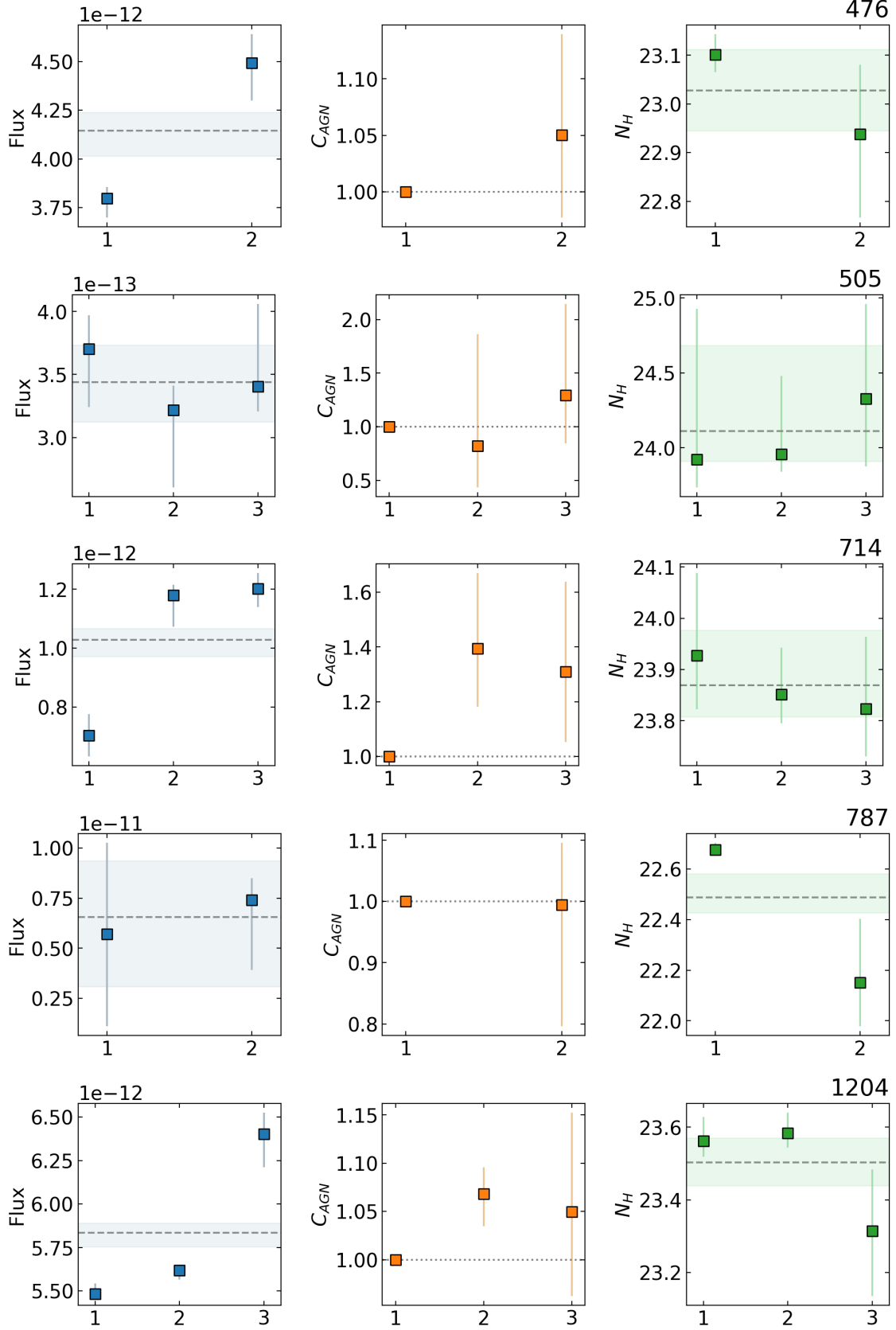


Figure 20. Same as Figure 19.

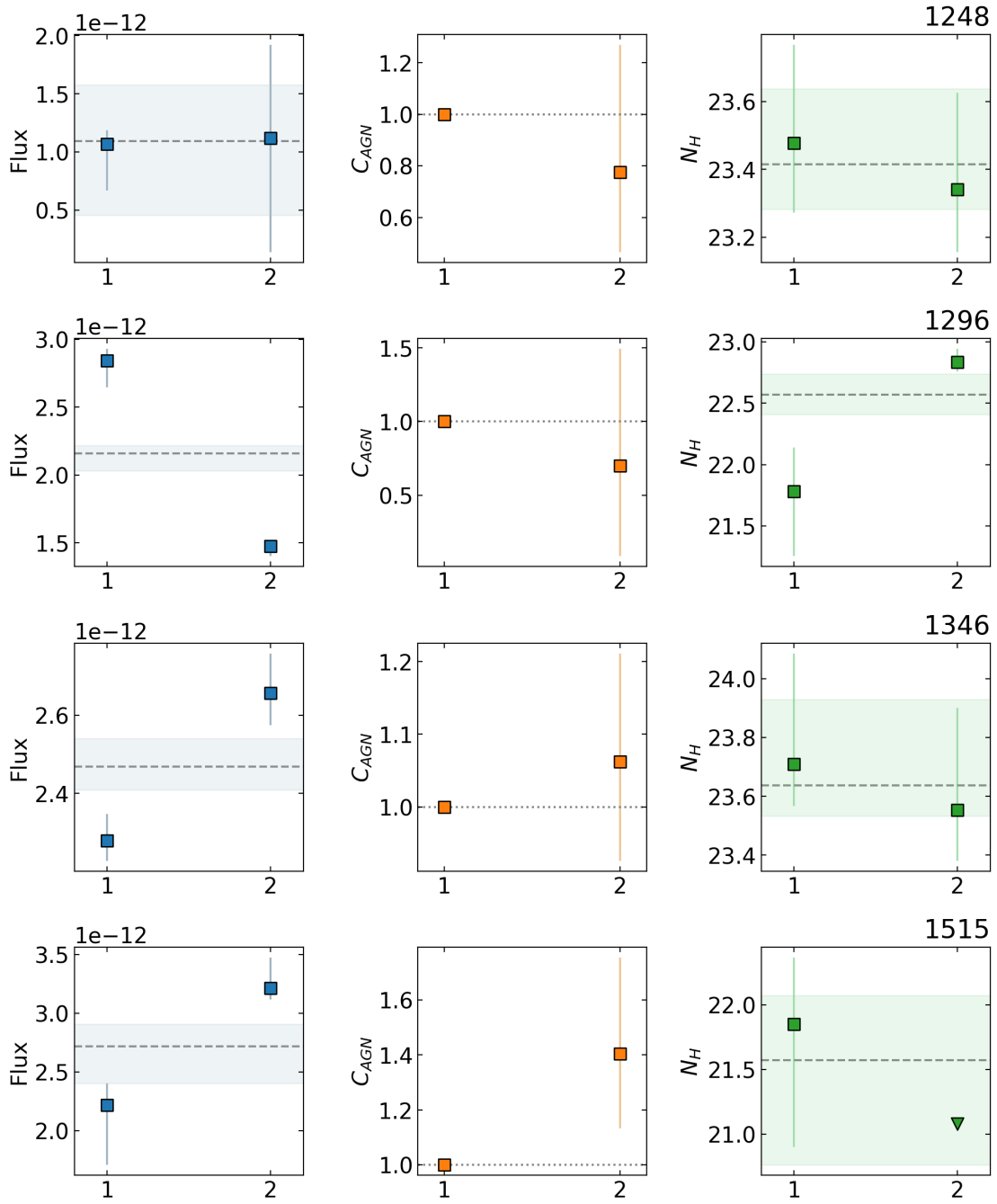


Figure 21. Same as Figure 19. For BAT ID 1515, the second observation yielded only an upper limit, which we indicate with a triangle, as the best-fit returned a value lower than the Galactic column density of $N_H = 1.2 \times 10^{21} \text{ cm}^{-2}$ (Kalberla et al. 2005).

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